



# Reverberation Mapping of Optical Emission Lines in Five Active Galaxies

M. M. Fausnaugh<sup>1</sup>, C. J. Grier<sup>2,3</sup>, M. C. Bentz<sup>4</sup>, K. D. Denney<sup>1,5,42</sup>, G. De Rosa<sup>6</sup>, B. M. Peterson<sup>1,5,6</sup>, C. S. Kochanek<sup>1,5</sup>, R. W. Pogge<sup>1,5</sup>, S. M. Adams<sup>7</sup>, A. J. Barth<sup>8</sup>, Thomas G. Beatty<sup>2,9</sup>, A. Bhattacharjee<sup>10,11</sup>, G. A. Borman<sup>12</sup>, T. A. Boroson<sup>13</sup>, M. C. Bottorff<sup>14</sup>, Jacob E. Brown<sup>15</sup>, Jonathan S. Brown<sup>1</sup>, M. S. Brotherton<sup>10</sup>, C. T. Coker<sup>1</sup>, S. M. Crawford<sup>16</sup>, K. V. Croxall<sup>17</sup>, Sarah Eftekharzadeh<sup>10</sup>, Michael Eracleous<sup>2,3,18</sup>, M. D. Joner<sup>19</sup>, C. B. Henderson<sup>20,43</sup>, T. W.-S. Holoien<sup>1,5</sup>, Keith Horne<sup>21</sup>, T. Hutchison<sup>14</sup>, Shai Kaspi<sup>22</sup>, S. Kim<sup>1</sup>, Anthea L. King<sup>23</sup>, Miao Li<sup>24</sup>, Cassandra Lochhaas<sup>1</sup>, Zhiyuan Ma<sup>15</sup>, F. MacInnis<sup>14</sup>, E. R. Manne-Nicholas<sup>4</sup>, M. Mason<sup>10</sup>, Carmen Montuori<sup>25</sup>, Ana Mosquera<sup>26</sup>, Dale Mudd<sup>1</sup>, R. Musso<sup>14</sup>, S. V. Nazarov<sup>12</sup>, M. L. Nguyen<sup>10</sup>, D. N. Okhmat<sup>12</sup>, Christopher A. Onken<sup>27</sup>, B. Ou-Yang<sup>4</sup>, A. Pancoast<sup>28,44</sup>, L. Pei<sup>9,29</sup>, Matthew T. Penny<sup>1,45</sup>, Radosław Poleski<sup>1</sup>, Stephen Rafter<sup>30</sup>, E. Romero-Colmenero<sup>16,31</sup>, Jessie Runnoe<sup>2,3,32</sup>, David J. Sand<sup>33</sup>, Jaderson S. Schimoia<sup>34</sup>, S. G. Sergeev<sup>12</sup>, B. J. Shappee<sup>35,46,47</sup>, Gregory V. Simonian<sup>1</sup>, Garrett Somers<sup>36,48</sup>, M. Spencer<sup>19</sup>, D. A. Starkey<sup>21</sup>, Daniel J. Stevens<sup>1</sup>, Jamie Tayar<sup>1</sup>, T. Treu<sup>37</sup>, Stefano Valenti<sup>38</sup>, J. Van Saders<sup>35,46</sup>, S. Villanueva Jr.<sup>1</sup>, C. Villforth<sup>39</sup>, Yaniv Weiss<sup>40</sup>, H. Winkler<sup>41</sup>, and W. Zhu<sup>1</sup>

<sup>1</sup> Department of Astronomy, The Ohio State University, 140 W. 18th Avenue, Columbus, OH 43210, USA

<sup>2</sup> Department of Astronomy & Astrophysics, The Pennsylvania State University, 525 Davey Laboratory, University Park, PA 16802, USA

<sup>3</sup> Institute for Gravitation and the Cosmos, The Pennsylvania State University, University Park, PA 16802, USA

<sup>4</sup> Department of Physics and Astronomy, Georgia State University, Atlanta, GA 30303, USA

<sup>5</sup> Center for Cosmology and AstroParticle Physics, The Ohio State University, 191 West Woodruff Avenue, Columbus, OH 43210, USA

<sup>6</sup> Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA

<sup>7</sup> Cahill Center for Astrophysics, California Institute of Technology, Pasadena, CA 91125, USA

<sup>8</sup> Department of Physics and Astronomy, 4129 Frederick Reines Hall, University of California, Irvine, CA 92697, USA

<sup>9</sup> Center for Exoplanets and Habitable Worlds, The Pennsylvania State University, 525 Davey Lab, University Park, PA 16802, USA

<sup>10</sup> Department of Physics and Astronomy, University of Wyoming, 1000 E. University Avenue, Laramie, WY, USA

<sup>11</sup> The Department of Biology, Geology, and Physical Sciences, Sul Ross State University, WSB 216, Box-64, Alpine, TX, 79832, USA

<sup>12</sup> Crimean Astrophysical Observatory, P/O Nauchny, Crimea 298409, Russia

<sup>13</sup> Las Cumbres Global Telescope Network, 6740 Cortona Drive, Suite 102, Santa Barbara, CA 93117, USA

<sup>14</sup> Fountainwood Observatory, Department of Physics FJS 149, Southwestern University, 1011 E. University Avenue, Georgetown, TX 78626, USA

<sup>15</sup> Department of Physics and Astronomy, University of Missouri, Columbia, USA

<sup>16</sup> South African Astronomical Observatory, P.O. Box 9, Observatory 7935, Cape Town, South Africa

<sup>17</sup> Illumination Works LLC., 5550 Blazer Parkway, Suite 152 Dublin, OH 43017, USA

<sup>18</sup> Center for Relativistic Astrophysics, Georgia Institute of Technology, Atlanta, GA 30332, USA

<sup>19</sup> Department of Physics and Astronomy, N283 ESC, Brigham Young University, Provo, UT 84602-4360, USA

<sup>20</sup> Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, CA 91109, USA

<sup>21</sup> SUPA Physics and Astronomy, University of St. Andrews, Fife, KY16 9SS, UK

<sup>22</sup> School of Physics and Astronomy, Raymond and Beverly Sackler Faculty of Exact Sciences, Tel Aviv University, Tel Aviv 69978, Israel

<sup>23</sup> School of Physics, University of Melbourne, Parkville, VIC 3010, Australia

<sup>24</sup> Department of Astronomy, Columbia University, 550 W. 120th Street, New York, NY 10027, USA

<sup>25</sup> DiSAT, Università dell'Insubria, via Valleggio 11, I-22100, Como, Italy

<sup>26</sup> Physics Department, United States Naval Academy, Annapolis, MD 21403, USA

<sup>27</sup> Research School of Astronomy and Astrophysics, Australian National University, Canberra, ACT 2611, Australia

<sup>28</sup> Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

<sup>29</sup> Department of Astronomy, University of Illinois at Urbana-Champaign, Urbana, IL 61801, USA

<sup>30</sup> Department of Physics, Faculty of Natural Sciences, University of Haifa, Haifa 31905, Israel

<sup>31</sup> Southern African Large Telescope Foundation, P.O. Box 9, Observatory 7935, Cape Town, South Africa

<sup>32</sup> Department of Astronomy, University of Michigan, 1085 S. University, 311 West Hall, Ann Arbor, MI 48109-1107, USA

<sup>33</sup> Physics & Astronomy Department, Texas Tech University, Box 41051, Lubbock, TX 79409-1051, USA

<sup>34</sup> Instituto de Física, Universidade Federal do Rio Grande do Sul, Campus do Vale, Porto Alegre, RS, Brazil

<sup>35</sup> Carnegie Observatories, 813 Santa Barbara Street, Pasadena, CA 91101, USA

<sup>36</sup> Department of Physics and Astronomy, Vanderbilt University, 6301 Stevenson Circle, Nashville, TN, 37235, USA

<sup>37</sup> Department of Physics and Astronomy, University of California, Los Angeles, CA 90095-1547, USA

<sup>38</sup> Department of Physics, University of California, One Shields Avenue, Davis, CA 95616, USA

<sup>39</sup> University of Bath, Department of Physics, Claverton Down, BA2 7AY, Bath, UK

<sup>40</sup> Physics Department, Technion, Haifa 32000, Israel

<sup>41</sup> Department of Physics, University of Johannesburg, PO Box 524, 2006 Auckland Park, South Africa

Received 2016 October 1; revised 2017 March 27; accepted 2017 April 3; published 2017 May 12

<sup>42</sup> NSF Postdoctoral Research Fellow.

<sup>43</sup> NASA Postdoctoral Program Fellow.

<sup>44</sup> Einstein Fellow.

<sup>45</sup> Sagan Fellow.

<sup>46</sup> Carnegie-Princeton Fellow.

<sup>47</sup> Hubble Fellow.

<sup>48</sup> VIDA Postdoctoral Fellow.



## Abstract

We present the first results from an optical reverberation mapping campaign executed in 2014 targeting the active galactic nuclei (AGNs) MCG+08-11-011, NGC 2617, NGC 4051, 3C 382, and Mrk 374. Our targets have diverse and interesting observational properties, including a “changing look” AGN and a broad-line radio galaxy. Based on continuum- $H\beta$  lags, we measure black hole masses for all five targets. We also obtain  $H\gamma$  and  $He\ II\ \lambda 4686$  lags for all objects except 3C 382. The  $He\ II\ \lambda 4686$  lags indicate radial stratification of the BLR, and the masses derived from different emission lines are in general agreement. The relative responsivities of these lines are also in qualitative agreement with photoionization models. These spectra have extremely high signal-to-noise ratios (100–300 per pixel) and there are excellent prospects for obtaining velocity-resolved reverberation signatures.

*Key words:* galaxies: active – galaxies: individual (MCG+08-11-011, NGC 2617, NGC 4051, 3C 382, Mrk 374)

*Supporting material:* machine-readable tables

## 1. Introduction

Understanding the interior structure of active galactic nuclei (AGNs) has been a major goal of extragalactic astrophysics since their identification as cosmological objects (Schmidt 1963). The current schematic structure of the central part of an AGN includes three main components: an accretion disk around a super-massive black hole (SMBH), a broad-line region (BLR), and an obscuring structure at some distance beyond the BLR. This basic picture accounts for the large luminosities and prominent recombination/excitation lines observed in Seyfert galaxy and quasar spectra (Burbidge 1967; Weedman 1977), as well as the dichotomy between Type 1 and Type 2 objects (Lawrence 1991; Antonucci 1993).

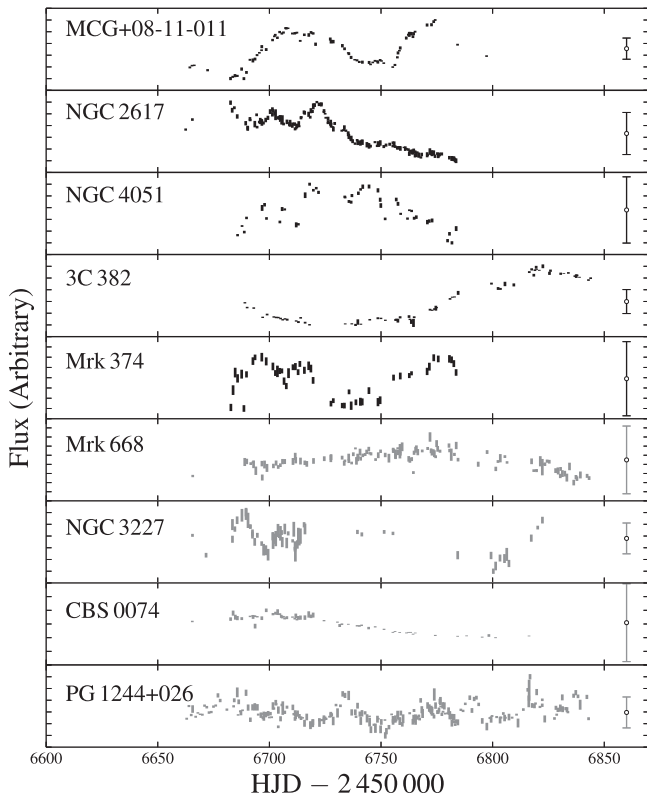
While this model has qualitatively explained the observational properties of AGNs, the details of the AGN interior structure remain poorly understood. The basic physics of the accretion disk are probably linked to the magnetorotational instability (Balbus & Hawley 1998), but it has not been possible to fully simulate an accretion disk and compare with observations (Koratkar & Blaes 1999; Yuan & Narayan 2014). It is also unclear if the BLR simply consists of ambient gas near the SMBH or if it is more directly connected with the accretion process. For example, broad-line-emitting gas might correspond to inflowing gas from large scales that feeds the accretion disk, or a portion of the BLR gas may be the result of an outflowing wind driven by radiation pressure from the accretion disk (Collin-Souffrin 1987; Murray & Chiang 1997; Elvis 2000; Proga & Kallman 2004; Proga & Kurosawa 2010; Higginbottom et al. 2014; Elitzur & Netzer 2016). The BLR could instead correspond to the portion of the obscuring structure lying within the dust sublimation radius (Netzer & Laor 1993; Simpson 2005; Gaskell et al. 2008; Nenkova et al. 2008; Mor & Netzer 2012). Other models explore the possibility that the accretion disk, BLR, and obscuring structure are not distinct at all, but are different observational aspects of a single structure bound to the central SMBH (e.g., Elitzur & Shlosman 2006; Czerny & Hryniewicz 2011; Goad et al. 2012).

Reverberation mapping (RM, Blandford & McKee 1982; Peterson 1993, 2014) is an effective way of investigating these scenarios. RM exploits the intrinsic variability of AGNs to investigate the matter distribution around the SMBH. The inner parts of the accretion disk emit in the far/extreme UV, providing ionizing photons that drive line emission from BLR gas. As the accretion disk stochastically varies, changes in the continuum flux are reprocessed as line emission by the BLR gas after a time delay that corresponds to the light-travel time across the BLR. Measuring this time delay (or “lag”) provides a

means of measuring the characteristic size scale of the line-emitting gas. Similarly, the UV continuum (or X-rays) deposits a small fraction of the accretion luminosity in the outer parts of the accretion disk and obscuring structure. Continuum variations will therefore change the local temperature of these structures, which can drive variable emission at longer continuum wavelengths—the outer part of the accretion disk emits primarily in the optical and the obscuring structure emits in the IR. By measuring any lag between the primary UV signal and light echoes at longer wavelengths, it is possible to “map” the size of the accretion disk and obscuring structure.

Early RM experiments were able to measure or constrain the physical scales of the three primary components: the accretion disk is on order a few light days from the SMBH (e.g., Wanders et al. 1997; Sergeev et al. 2005), the BLR ranges from several light days to a few light months or light years, depending on the AGN luminosity (Wandel et al. 1999; Kaspi et al. 2000; Peterson et al. 2004; Kaspi et al. 2005), and the obscuring structure extends several light months or light years beyond the BLR (Clavel et al. 1989; Oknyanskij & Horne 2001; Suganuma et al. 2006). More recent RM studies have provided additional details. The detection of continuum lags across the accretion disk provides information about the disk’s temperature gradient, and it appears that the disks are somewhat larger than the predictions from standard models (e.g., Shappee et al. 2014; Edelson et al. 2015; Fausnaugh et al. 2016; McHardy et al. 2016), as also found in microlensing studies of lensed quasars (e.g., Morgan et al. 2010; Blackburne et al. 2011; Mosquera et al. 2013). Mid- to far-IR echoes from the obscuring structure have facilitated investigation of AGN dust properties and suggest that the obscuring structure is clumpy and has a mixed chemical composition (Kishimoto et al. 2007; Vazquez et al. 2015).

RM of the BLR is of particular importance for AGN studies because velocity information in the broad-line profile combined with the observed time delay provides a well-calibrated estimate of the SMBH mass. Approximately 60 AGNs have RM mass measurements (Bentz & Katz 2015), and this sample anchors the scaling relations used to infer the majority of SMBH masses throughout the universe (e.g., McLure & Dunlop 2004; Vestergaard & Peterson 2006; Trakhtenbrot & Netzer 2012; Park et al. 2013; Mejía-Restrepo et al. 2016, and references therein). New insights into the BLR structure have also become available with velocity-resolved analyses (e.g., Bentz et al. 2010; Denney et al. 2010; Barth et al. 2015; Valenti et al. 2015; Du et al. 2016). By combining information about the BLR time delay as a function of line of sight velocity, it is possible to distinguish among geometric and dynamical



**Figure 1.** g-band light curves of all targets in the 2014 monitoring campaign except for Akn 120 and NGC 5548 (see Section 2.1). In this study, we focus on MCG+08-11-011, NGC 2617, NGC 4051, 3C 382, and Mrk 374. The extent of the error bars on the open white circles represent 10% variations for each flux scale.

configurations, such as flattened versus spherical matter distributions and dynamics dominated by rotation, infall, or outflow (Horne 1994; Horne et al. 2004; Bentz et al. 2010; Grier et al. 2013b; Pancoast et al. 2014a, 2014b). So far, only about 10 AGNs have such detailed velocity-resolved results, but they suggest a wide range of dynamics and geometries.

In this work, we present the first results from an intensive RM campaign executed in 2014. This campaign had two primary goals: to measure SMBH masses in several objects with interesting or peculiar observational properties, and to expand the sample of AGNs with velocity-resolved reverberation signatures. NGC 5548 was also observed in this campaign as part of the multiwavelength AGN STORM project (De Rosa et al. 2015; Edelson et al. 2015; Fausnaugh et al. 2016; Goad et al. 2016). Ground-based spectroscopic results for this object are presented by Pei et al. (2017). Here, we present the final data and initial analysis of other AGNs from this campaign, reporting continuum and line light curves, continuum-line lag measurements, and SMBH masses for five objects. We detected variability in the  $H\beta$ ,  $H\gamma$ , and  $He\ II\ \lambda 4686$  emission lines for most objects, which we also use to explore the photoionization conditions in the BLR (Korista & Goad 2004; Bentz et al. 2010). These data are of exceptional quality and should allow us recover velocity-resolved reverberation signatures in future work.

In Section 2, we present our target AGNs, observations, data reduction, and light curves. In Section 3, we explain our time-series analysis and report continuum-line lags. In Section 4, we measure the gas velocities and estimate SMBH masses. In Section 5, we discuss our results, and in Section 6, we

summarize our findings. We assume a consensus cosmology with  $H_0 = 70\text{ km s}^{-1}\text{ Mpc}^{-1}$ ,  $\Omega_m = 0.3$ , and  $\Omega_\Lambda = 0.7$ .

## 2. Observations and Data Reduction

### 2.1. Targets

In spring of 2014 we monitored 11 AGNs over the course of a six-month RM campaign. The AGNs were selected with the aim of expanding the database of RM SMBH masses, particularly for objects with diverse and peculiar observational characteristics. The second goal of our campaign was to investigate the dynamics and geometry of the BLR with velocity-resolved reverberation signatures, i.e., velocity-delay maps and dynamical models (see, e.g., Grier et al. 2013b; Pancoast et al. 2014a). Here, we focus on results related to SMBH masses and we will pursue the velocity-resolved analysis in future work.

Figure 1 shows g-band light curves from the Las Cumbres Observatory (LCO) 1 m network for nine of our targets (we discuss these data in detail in Section 2.3). Not shown are Akn 120, which was dropped early in the campaign because of low variability, and NGC 5548, for which the results are presented elsewhere (Fausnaugh et al. 2016; Pei et al. 2017). In order to estimate a black hole mass, we must measure a continuum-line lag. We have not been able to measure such a reverberation signal for Mrk 668, NGC 3227, CBS 0074, and PG 1244+026. These sources have lower signal-to-noise ratios (S/Ns) than the other objects (generally 30–70 per pixel, although NGC 3227 was  $\sim 90$  per pixel; see Section 2.5.3), and they display lower variability amplitudes. The fractional root-mean-square amplitude ( $F_{\text{var}}$  as defined in Section 2.5.3 below) is 0.012 for Mrk 668, 0.037 for NGC 3227, 0.010 for CBS 0074, and 0.025 for PG 1244+026. For Mrk 668, the slow rate of change in the light curve also makes it impossible to measure short lags. For NGC 3227, the light curve is problematic because of the limited sampling and large gaps; however, this object was also observed during a monitoring campaign in 2012, and we will combine the data from both campaigns in a future analysis. For CBS 0074 and PG 1244+026, we have not been able to obtain a sufficiently precise calibration of the spectra (see Section 2.2.2) to detect emission-line variability.

We succeeded in measuring black hole masses for MCG+08-11-011, NGC 2617, NGC 4051, 3C 382, and Mrk 374. Table 1 lists the some of the important properties of these objects (several of which are measured in this study), and we provide additional comments as follows.

- i. MCG+08-11-011 is a strong X-ray source for which spectral signatures of a relativistically broadened Fe  $K\alpha$  line have been observed with *Suzaku* (Bianchi et al. 2010). The Fe  $K\alpha$  emission is believed to be emitted close to the inner edge of the accretion disk and can potentially be used to measure the spin parameter of the black hole. Because the black hole mass and spin are to some extent degenerate when fitting the broad Fe  $K\alpha$  profile, an independent mass estimate from RM can greatly assist with the spin measurement.
- ii. NGC 2617 was discovered by Shappee et al. (2014) to be a “changing look” AGN. In 2013, after a large X-ray/optical outburst, follow-up spectroscopic observations showed the presence of broad lines, while archival spectra from 2003 show only a weak broad component of



**Table 1**  
Target Properties

Object	Redshift	$D_L$ (Mpc)	Number of Good-weather Epochs	$F$ [O III] $\lambda$ 5007 ( $10^{-14}$ erg cm $^{-2}$ s $^{-1}$ )	[O III] $\lambda$ 5007 Light Curve Scatter (%)	$\log \lambda L_{5100 \text{ \AA}}$ (erg s $^{-1}$ )	$\log \lambda L_{\text{host}}$ (erg s $^{-1}$ )	$E(B-V)$ (mag)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
MCG+08-11-011	0.0205	89.1	6	$61.33 \pm 0.21$	0.09	43.59	43.28	0.19
NGC 2617	0.0142	61.5	3	$6.88 \pm 0.06$	1.37	43.12	42.95	0.03
NGC 4051	0.0023	17.1	3	$41.30 \pm 0.26$	0.36	42.38	42.23	0.01
3C 382	0.0579	258.7	6	$7.87 \pm 0.06$	0.92	44.20	43.98	0.06
Mrk 374	0.0426	188.5	3	$7.27 \pm 0.11$	0.62	43.98	43.61	0.05

**Note.** Column 2 is taken from the NASA Extragalactic Database. Column 3 gives the luminosity distance  $D_L$  in a consensus cosmology, except for NGC 4051 for which the luminosity distance is from Tully et al. (2008). Column 4 gives the number of nights with clear and stable conditions on which each object was observed. Each object had three observations per night, which were used to calculate the narrow [O III] $\lambda$ 5007 line flux. The line flux and its uncertainty are given in Column 5. Column 6 gives the fractional variation of the [O III] $\lambda$ 5007 line light curve, which serves as an estimate of the night-to-night calibration error (Section 2.5.1). Column 7 gives the observed luminosity (corrected for Galactic extinction), calculated from the observed 5100 Å rest-frame light curve and Column 3. Column 8 gives the luminosity of the host-galaxy starlight in the spectroscopic extraction aperture, also corrected for Galactic extinction (Section 5.1). Note that Column 7 includes the contribution from the host galaxy. Column 9 gives the Galactic reddening value from Schlafly & Finkbeiner (2011).

H $\alpha$ . This means that the classification of NGC 2617 changed from a Seyfert 1.9 to Seyfert 1.0 sometime in the intervening decade. Few optical “changing look” AGNs are known, although systematic searches through long-term survey data (such as the SDSS, LaMassa et al. 2015; MacLeod et al. 2016) and targeted repeat spectroscopy (Ruan et al. 2016; Runco et al. 2016; Runnoe et al. 2016) have recently expanded the sample size to approximately 20 objects, depending on how “changing look” AGNs are defined. The absolute rate of this phenomenon is very uncertain, but these recent studies suggest that it may be relatively common over several decades, a timescale that long-term spectroscopic surveys are only beginning to probe. Velocity-resolved dynamical information is of special interest in an object such as this, since the presence of outflows or infall may provide clues about the physical mechanism behind the change in Seyfert category.

- iii. NGC 4051 has been the target of several optical and X-ray RM campaigns (Peterson et al. 2000, 2004; Shemmer et al. 2003; Denney et al. 2009b; Miller et al. 2010; Turner et al. 2016). However, the short H $\beta$  lag, comparable to the cadence of most monitoring campaigns, has led to mixed and inconsistent results. Denney et al. (2009b) found an H $\beta$  lag of  $1.87 \pm 0.52$  days, roughly a factor of two smaller than previous studies. Because of the large change, as well as the lag’s small value compared to the monitoring cadence, we re-observed NGC 4051 during the 2014 campaign to check this result. For one month of the campaign (2014 February 17 to March 16 UTC), we also increased the monitoring cadence of NGC 4051 to twice nightly, in order to securely resolve the expected short H $\beta$  lag.

NGC 4051 is also an archetypal narrow-line Seyfert 1 (NLS1), meaning that the width of its H $\beta$  line is  $\lesssim 2000$  km s $^{-1}$ . There are two competing theories to explain the NLS1 phenomenon: high accretion rates or rotationally dominated BLR dynamics seen nearly face-on. Both explanations can account for the narrow linewidths given the AGN luminosity. Insight into the structure of the BLR can help distinguish between these explanations, so there is considerable interest in reconstructing a velocity-delay map for this object.

- iv. 3C 382 is an FR II broad-line radio galaxy (Osterbrock et al. 1975, 1976). Few radio-loud AGNs have RM mass measurements, although there are notable examples such as 3C 390 (Shapovalova et al. 2010; Dietrich et al. 2012), 3C 273 (Kaspi et al. 2000; Peterson et al. 2004), and 3C 120 (Peterson et al. 2004; Grier et al. 2012). These objects are typically more luminous than radio-quiet AGNs, so they have large lags (on the order months to years) that are difficult and expensive to measure. However, radio emission is thought to be associated with more massive black holes, which can be tested by anchoring radio-loud AGNs to the RM mass scale. Radio jets can also provide an indirect estimate of the inclination of the BLR, if the BLR is a disk structure with the rotation axis aligned to that of the jet (Wills & Browne 1986). Several jet-orientation indicators exist for 3C 382, and Eracleous et al. (1995) estimated the BLR inclination in 3C 382 using dynamical models of the double-peaked H $\alpha$  profile. Velocity-delay maps and dynamical models would provide an interesting comparison to these estimates.
- v. We observed Mrk 374 in an RM campaign from 2012, but this AGN did not display sufficient variability to measure emission-line lags at that time. Although Mrk 374 is our least variable source, we succeeded in measuring a line lag from the 2014 campaign, and we present the first RM-based black hole mass here.

## 2.2. Spectra

### 2.2.1. Observations

We obtained spectra on an approximately daily cadence between 2014 January 04 and July 06 UTC using the Boller and Chivens CCD Spectrograph on the 1.3 m McGraw-Hill telescope at the MDM Observatory. We used the 350 mm $^{-1}$  grating, yielding a dispersion of 1.33 Å per pixel with wavelength coverage from 4300 to 5600 Å. We kept the position angle of the slit fixed to 0° for the entire campaign, with a slit width of 5''0 to minimize losses due to differential refraction and aperture effects caused by extended emission (i.e., the host-galaxy and narrow-line region, Peterson et al. 1995). Because of the large slit width, the spectroscopic resolution for point sources (such as the AGN) is limited by the

**Table 2**  
Observed-frame Integration Windows

Object	5100 Å (Å)	H $\beta$ (Å)	H $\gamma$ (Å)	He II $\lambda$ 4686 (Å)	[O III] $\lambda$ 5007 (Å)	[O III] $\lambda$ 4959 (Å)
MCG+08-11-011	5190–5230	4890–5040	4375–4485	4650–4890	5085–5130	5040–5075
NGC 2617	5170–5200	4835–5050	4310–4520	4620–4835	5055–5093	5010–5040
NGC 4051	5115–5145	4835–4920	4315–4390	4610–4740	5000–5045	4955–4977
3C 382	5380–5400	5000–5270	4425–4745	4795–4930	5275–5330	5228–5257
Mrk 374	5315–5350	4995–5140	4490–4580	4765–4915	5205–5245	5160–5187

image seeing. We discuss this in more detail in Section 4, but comparison with high-resolution observations suggest that the effective spectral resolution is approximately 7.0 Å.

The two-dimensional spectra were reduced using standard IRAF tasks for overscan, bias, and flat-field corrections, and cosmic rays were removed using LA-cosmic (van Dokkum 2001). We extracted one-dimensional spectra from a 15''0 window centered on a linear fit to the trace, and we derived wavelength solutions from comparison lamps taken in the evening and morning of all observing nights. We also corrected for zero-point shifts in the wavelength solutions (due to flexure in the telescope) by taking xenon lamp exposures just prior to each observing sequence. However, every AGN was observed for a series of three 20 minute exposures and the wavelength zero-point can drift over the course of this hour, especially at high airmass. We therefore tie the wavelength solution of the first exposure to the contemporaneous xenon lamp, and then apply shifts that align the [O III]  $\lambda$ 5007 emission line of subsequent exposures to that of the first. This procedure results in wavelength solutions accurate to 0.56 Å, as measured from night-sky emission lines.

We applied relative flux calibrations using sensitivity curves derived from nightly observations of standard stars. For most of the campaign, we use Feige 34 (Oke 1990) to define the nightly sensitivity curve. However, this star began to set near dusk at the end of the campaign, so we tied our relative flux calibration to BD+33°2642 (Oke 1990) for the final two weeks. The change in standard star could potentially result in a systematic change in the observed continuum slopes. However, BD+33°2642 and Feige 34 were observed for a one-month overlap period before the transition, and the sensitivity curves derived from both stars agree well during this time period. Of the targets presented here, only 3C 382 was observed during the final two weeks, and we did not find any anomalous changes in the spectral slope during this period. As a check on the relative flux calibration, we also looked for a “bluer when brighter” trend, caused by an increasing fraction of host-galaxy light when the AGN is in a faint state and/or intrinsic variations in the AGN spectral energy distribution (e.g., Wilhite et al. 2005; Sakata et al. 2010). We measured the spectral slope by fitting a straight line to each spectrum with the emission lines masked, and for all cases except the weakly varying Mrk 374, we found a significant anticorrelation between the mean flux and the spectral slope. Detecting the “bluer when brighter” effect lends additional confidence to our relative flux calibration.

We also obtained six epochs of observations with the 2.3 m telescope at Wyoming Infrared Observatory (WIRO) and the WIRO Long Slit Spectrograph. The WIRO spectra were used to fill in gaps in the MDM monitoring, and we matched the spectrograph configuration to that of the MDM spectrograph as closely as possible. This includes a 5''0 slit at position angle 0° for all observations, and we used the same extraction/sky

apertures as for the MDM observations. The wavelength calibrations and spectral slopes of the WIRO data agree well with the MDM observations, and we discuss the calibration of the WIRO data to the MDM flux scale in Section 2.5.1.

### 2.2.2. Night-to-night Flux Calibration

In order to account for variable atmospheric extinction and seeing, we employ the calibration algorithms introduced by Fausnaugh (2017). This approach is similar to the older method of van Groningen & Wanders (1992), but yields markedly better calibrations. We assume that the [O III]  $\lambda$ 5007 emission line is constant over the course of our campaign and we transform the observed spectra so that their [O III]  $\lambda$ 5007 line profiles match those of the “photometric” nights (nights with clear conditions and stable seeing). We treat the WIRO and MDM spectra separately and inter-calibrate the two flux scales below (Section 2.5.1).

Fausnaugh (2017) discussed the details of our implementation and provided a python package (`mapspec`<sup>49</sup>) to build and apply a rescaling model to time-series spectra. For completeness, we briefly outline the procedure here.

- i. First, we collected the spectra taken on photometric nights (as judged by the observers on site) and estimated their [O III]  $\lambda$ 5007 line fluxes. The line fluxes were measured by subtracting a linearly interpolated estimate of the local continuum underneath the line and then integrating the remaining flux using Simpson’s method. We provide the wavelength regions of the integration and the continuum fit in Tables 2 and 3. We applied iterative  $3\sigma$  clippings to the line fluxes, where  $\sigma$  is their root-mean-square (rms) scatter, in order to reject any outliers (due to slit losses or anomalies in the sky conditions). We then averaged the remaining flux measurements to estimate the true line flux. The measured [O III]  $\lambda$ 5007 line fluxes for each object are given in Table 1. Table 1 also gives the number of photometric epochs used to determine these fluxes for each AGN (we usually took three spectra per epoch).
- ii. We then combined the remaining photometric spectra into a reference spectrum using a noise-weighted average. In this step, any residual wavelength shifts were removed by aligning the [O III]  $\lambda$ 5007 line profiles using Markov Chain Monte Carlo (MCMC) methods—the spectra are shifted by the wavelength shift that minimizes the sum of the squares of residuals between the [O III]  $\lambda$ 5007 line profiles. Linear interpolation is used for wavelength shifts of fractional pixels.
- iii. Due to changes in seeing, spectrograph focus, and small guiding errors, the spectral resolution of each observation

<sup>49</sup> <https://github.com/mmfausnaugh/mapspec>

**Table 3**  
Observed-frame Continuum-fitting Windows

Object	Line Side	H $\beta$ (Å)	H $\gamma$ (Å)	He II $\lambda$ 4686 (Å)	[O III] $\lambda$ 5007 (Å)	[O III] $\lambda$ 4959 (Å)
MCG+08-11-011	Blue	4860–4890	4360–4375	4620–4650	5075–5085	5030–5040
	Red	5130–5150	4485–4500	4860–4880	5130–5150	5075–5085
NGC 2617	Blue	4820–4835	4300–4310	4585–4620	5050–5055	5000–5010
	Red	5110–5150	4520–4535	4820–4835	5093–5098	5040–5050
NGC 4051	Blue	4800–4835	4300–4315	4605–4615	4990–5000	4945–4955
	Red	4920–4950	4390–4400	4740–4775	5045–5055	4977–4990
3C 382	Blue	4975–5000	4410–4425	4785–4795	5265–5275	5218–5228
	Red	5385–5425	4745–4760	4930–4940	5330–5340	5257–5271
Mrk 374	Blue	4970–4990	4455–4490	4690–4765	5190–5205	5150–5160
	Red	5140–5160	4580–4600	4915–5000	5245–5255	5187–5195

is slightly different. To address this, we smooth the reference spectrum with a Gaussian kernel so that the [O III]  $\lambda$ 5007 linewidth matches the largest [O III]  $\lambda$ 5007 linewidth in the time series. The smoothed reference spectrum will define the final resolution of the calibrated spectra.

- iv. The time-series spectra are then aligned to the reference by matching the [O III]  $\lambda$ 5007 line profiles, again in a least-squares sense using MCMC methods. The differences in line profiles are modeled by a flux rescaling factor, a wavelength shift, and a smoothing kernel. After rescaling, we combine spectra from a single night using a noise-weighted average.

### 2.3. Imaging

Our spectroscopic observations are supplemented with broadband imaging observations. Contributing telescopes were the 0.7 m at the Crimean Astrophysical Observatory (CrAO), the 0.5 m Centurian 18 at Wise Observatory (WC18, Brosch et al. 2008), and the 0.9 m at West Mountain Observatory (WMO). CrAO uses an AP7p CCD with a pixel scale of  $1''.76$  and a  $15' \times 15'$  field of view, WC18 uses a STL-6303E CCD with a pixel scale of  $1''.47$  and a  $75' \times 50'$  field of view, and WMO uses a Finger Lakes PL-3041-UV CCD with a pixel scale of  $0''.61$  and a field of view of  $21' \times 21'$ . Fountainwood Observatory (FWO) also provided observations of NGC 4051 with a 0.4 m telescope using an SBIG 8300M CCD. The pixel scale of this detector is  $0''.35$  and the field of view is  $19' \times 15'$ . All observations were taken with the Bessell *V*-band.

In addition, we obtained *ugriz* imaging with the LCO 1 m network (Brown et al. 2013), which consists of nine identical 1 m telescopes at four observatories spread around the globe. These data were originally acquired as part of LCO’s AGN Key project (Valenti et al. 2015). The main goal is to search for continuum reverberation signals, which we will pursue in a separate study (Fausnaugh et al. 2017). However, 3C 382 and Mrk 374, which are our faintest sources, had low variability amplitudes and poorer S/Ns, so we included the LCO *g*-band data in the continuum light curves of these objects. Each LCO telescope has the same optic system and detectors—at the time of the RM campaign, the detectors were SBIGSTX-16803

cameras with a field of view of  $16' \times 16'$  and a pixel scale of  $0''.23$ .

We analyzed the imaging data using the image subtraction software (ISIS) developed by Alard & Lupton (1998). Images were first uploaded to a central repository and vetted by eye for obvious reduction errors or poor observing conditions. We then registered the images to a common coordinate system and constructed a high S/N reference frame by combining the best-seeing and lowest-background images. When combining, ISIS adjusts the images to a common seeing by convolving the point-spread function (PSF) of each image with a spatially variable kernel. Finally, we subtracted the reference frame from each image, again allowing ISIS to match the PSFs using its convolution routine. Reference images and subtractions for each telescope/filter/detection system were constructed separately—we discuss combining the photometric measurements in Section 2.5.2.

### 2.4. Mean and rms Spectra

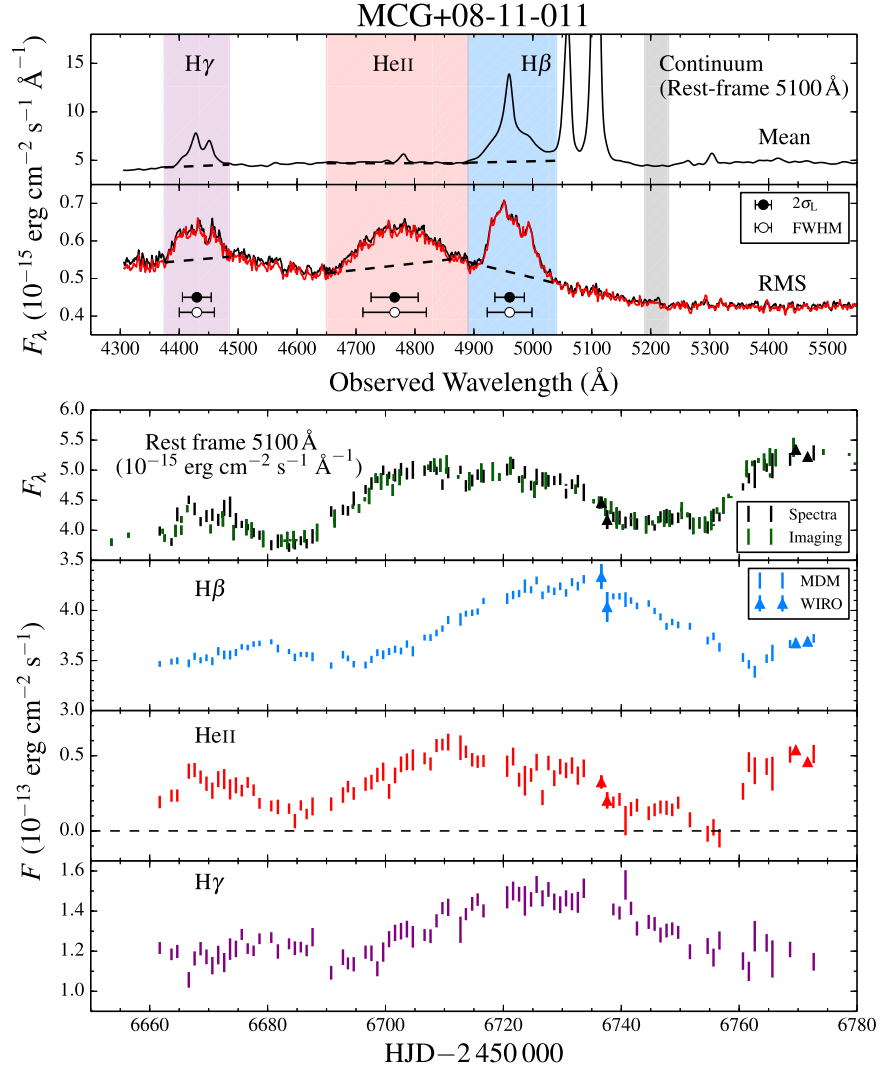
Figures 2–6 show the noise-weighted mean spectrum

$$\bar{F}(\lambda) = \frac{\sum_{i=1}^{N_t} F(\lambda, t_i) / \sigma^2(\lambda, t_i)}{\sum_{i=1}^{N_t} 1 / \sigma^2(\lambda, t_i)} \quad (1)$$

for each object using the MDM observations, where  $F(\lambda, t_i)$  is the flux density at epoch  $t_i$  and  $\sigma(\lambda, t_i)$  is its uncertainty. Figures 2–6 also show root-mean-square (rms) residual spectra, defined as

$$\sigma_{\text{rms}}(\lambda) = \sqrt{\frac{1}{N_t - 1} \sum_{i=1}^{N_t} [F(\lambda, t_i) - \bar{F}(\lambda)]^2}. \quad (2)$$

By the Wiener–Khinchin theorem, this statistic is proportional to the integrated variability power at each wavelength, so  $\sigma_{\text{rms}}$  is free of constant contaminants such as host-galaxy and narrow emission-line flux. However, the total variability power contains contributions from both intrinsic variations and from statistical fluctuations/measurement uncertainties. In order to separate these components, we use a maximum-likelihood method (see Park et al. 2012b; Barth et al. 2015; De Rosa et al. 2015). We solve for the intrinsic variability  $\sigma_{\text{var}}(\lambda)$  that



**Figure 2.** Summary of the data for MCG+08-11-011. Top panel: mean spectrum of the full time series (top) and rms spectrum (bottom), as defined in Section 2.4. The black line is  $\sigma_{\text{rms}}(\lambda)$  (Equation (2)), and the red line is  $\sigma_{\text{var}}(\lambda)$ , which includes a correction for measurement uncertainties (Equation (3)). The shaded regions show the windows from which continuum and line light curves were extracted. The dashed lines show local linear fits to the continuum underlying the lines. The error bars show the rms linewidth ( $\sigma_L$ ) and full-width at half maximum (FWHM). Bottom panel: light curves for the rest-frame 5100 Å continuum (imaging data is shown in green) and optical recombination lines. Over-subtraction of the continuum occasionally results in negative He II  $\lambda 4686$  fluxes, although RM measurements are only sensitive to the relative variations.

minimizes the negative log-likelihood

$$-2 \ln \mathcal{L} = \sum_{i=1}^{N_t} \frac{[F(\lambda, t_i) - \hat{F}(\lambda)]^2}{\sigma^2(\lambda, t_i) + \sigma_{\text{var}}^2(\lambda)} + \sum_{i=1}^{N_t} \ln [\sigma^2(\lambda, t_i) + \sigma_{\text{var}}^2(\lambda)], \quad (3)$$

where  $\hat{F}(\lambda)$  is the “optimal average” weighted by  $\sigma^2(t_i) + \sigma_{\text{var}}^2$ . We self-consistently fit for  $\hat{F}(\lambda)$  while solving for  $\sigma_{\text{var}}(\lambda)$ , and we show the estimate of  $\sigma_{\text{var}}(\lambda)$  with the red lines in Figures 2–6. In the limit that  $\sigma(\lambda, t_i) \rightarrow 0$ , it is clear that  $\sigma_{\text{var}}$  is equivalent to  $\sigma_{\text{rms}}$ . For high S/N data such as these,  $\sigma_{\text{var}}(\lambda)$  is nearly equal to  $[\sigma_{\text{rms}}^2(\lambda) - \bar{\sigma}^2(\lambda)]^{1/2}$ , where  $\bar{\sigma}^2(\lambda)$  is the average of the squared measurement uncertainties across

the time series:

$$\bar{\sigma}^2(\lambda) = \frac{1}{N} \sum_{i=1}^{N_t} \sigma^2(\lambda, t_i). \quad (4)$$

The overall effect is to reduce the squared amplitude of the variability spectrum by the mean squared measurement uncertainty—in all objects except for Mrk 374, this effect is negligible.

## 2.5. Light Curves

### 2.5.1. Spectroscopic Light Curves

We extracted spectroscopic light curves for the wavelength windows listed in Table 2 for each AGN. We chose these windows based on visual inspection of the variable line profiles in the  $\sigma_{\text{var}}(\lambda)$  spectra, with the main goal of capturing the

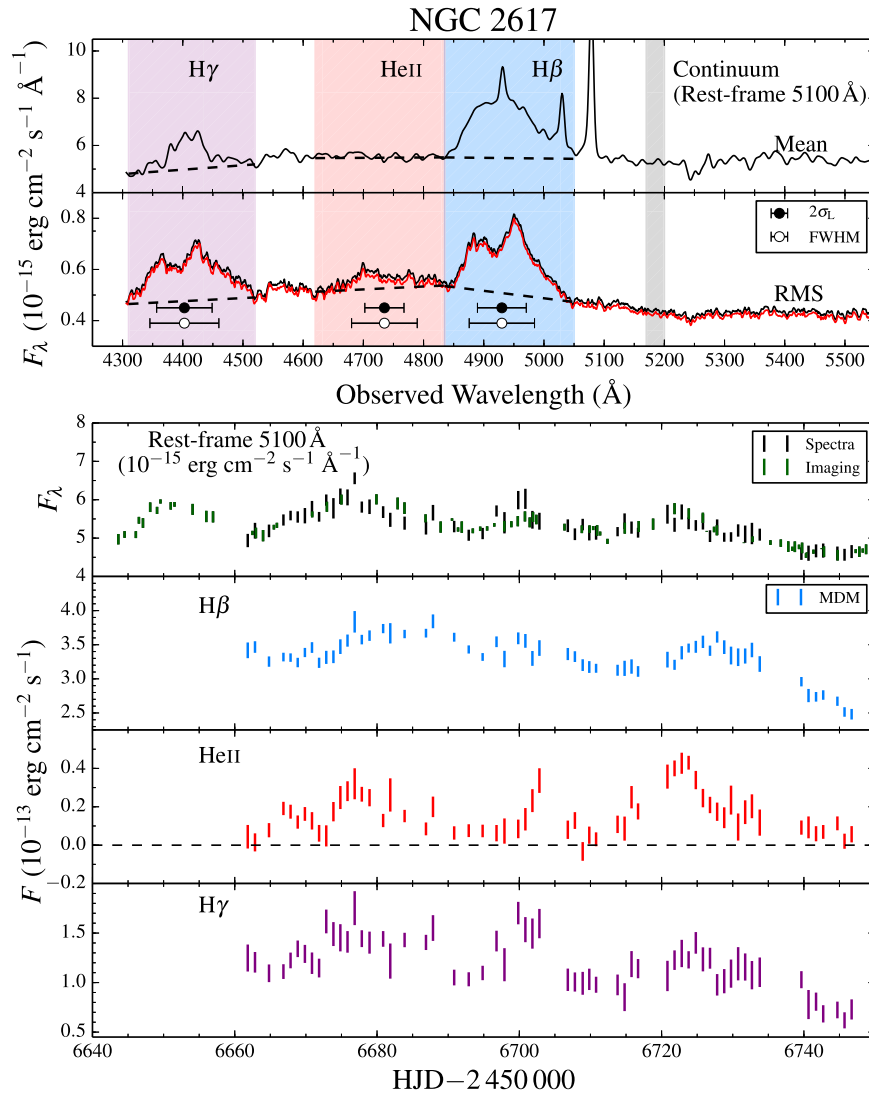


Figure 3. Same as Figure 2 but for NGC 2617.

strongest variations in the lines. For 3C 382, the component tentatively identified as He II  $\lambda 4686$  is blueshifted by almost 100 Å relative to the systematic redshift, and if variable He II  $\lambda 4686$  has a similar profile as the Balmer lines in this object, this component corresponds to the blue wing of the line.

The rest-frame 5100 Å continuum, which is relatively free of emission/absorption lines, was estimated by averaging the flux density in the listed wavelength region. Emission-line fluxes were determined in the same way as for the [O III]  $\lambda 5007$  line. First, we subtracted a linear least-squares fit to the local continuum underneath the emission line. Wavelength regions for the continuum fits are given in Table 3. Then, we integrated the remaining flux using Simpson’s method (we did not assume a functional form for the emission line). In cases where the broad H $\beta$  wing extends underneath [O III]  $\lambda 4959$ , we subtracted the narrow emission line (again with a local linear approximation of the underlying flux) before integrating the broad line. We did not attempt to separate the narrow components of H $\beta$  and H $\gamma$  from the broad components. These narrow components act as constant flux-offsets for the light curves.

The continuum estimates can lead to significant systematic uncertainties because the continuum-fitting windows may be contaminated by broad-line wing emission, and the local

linearly interpolated continuum may leave residual continuum flux to be included in the line profile. Both of these effects can introduce spurious correlations between the continuum and line light curves, which may biased the final lag estimates. Because we use the  $\sigma_{\text{var}}(\lambda)$  spectra to select the line and continuum windows, variability in the line wings probably does not have a large impact on our results, and we have found the the resulting light curves (and their lags) are robust to 5 to 10 angstrom changes in the continuum and line windows. Larger shifts, especially as the continuum-fitting windows move further from the lines, can result in significantly different lags (on the order of three times the statistical uncertainties). Full spectral decompositions may be able to address this issue in future studies (see Barth et al. 2015 for a detailed discussion). We discuss these systematic uncertainties further in Section 4.

After we extracted line fluxes from the WIRO and MDM spectra, we combined the measurements by forcing the light curves to be on the same flux scale. We used the mean MDM [O III]  $\lambda 5007$  line to define this scale and multiplied the WIRO line fluxes so that the mean value matched that of MDM. A more sophisticated inter-calibration model would include an additive offset, to account for different amounts of host-galaxy starlight in the MDM and WIRO spectra. However, with the



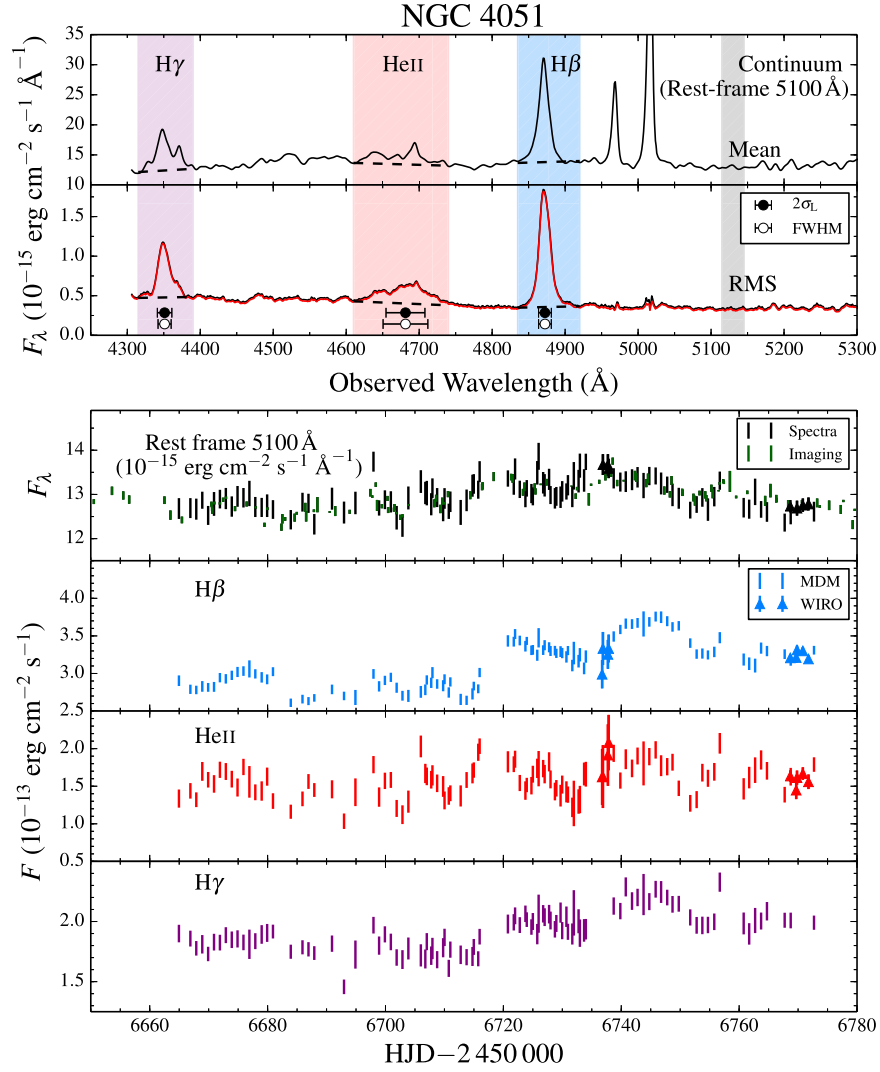


Figure 4. Same as Figure 2 but for NGC 4051.

limited amount of WIRO data, additional calibration parameters cannot be well-constrained, and we found the simple multiplicative approach to be adequate. The required rescaling factors were 1.21 for MCG+08-11-011, 1.14 for NGC 4051, 1.09 for 3C 382, and 1.73 for Mrk 374. Weather at WIRO prevented observations of NGC 2617.

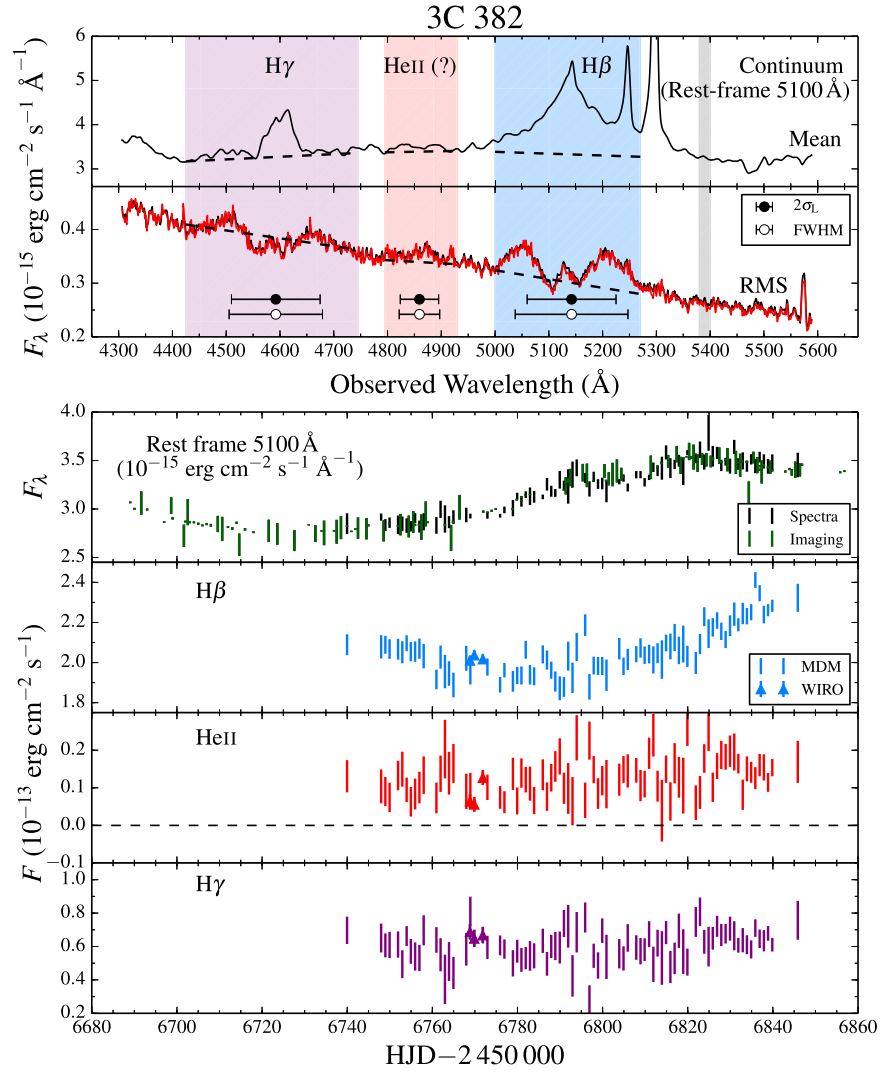
The statistical uncertainty on the continuum flux was estimated from the standard deviation within the wavelength region,

$$\sigma(t_j) = \sqrt{\frac{1}{N_\lambda - 1} \sum_{i=1}^{N_\lambda} [F(\lambda_i, t_j) - \bar{F}(t_j)]^2}, \quad (5)$$

where  $\bar{F}(t_j)$  is the evenly weighted average flux density at epoch  $t_j$ . Uncertainties on the line light curves were estimated using a Monte Carlo approach: we perturbed the observed spectrum with random deviates scaled to the uncertainty at each wavelength, subtracted a new estimate of the underlying continuum (and the narrow [O III]  $\lambda 4959$  line when appropriate), and re-integrated the line flux. The deviates were drawn from the multivariate normal distribution defined by the covariance matrix of the rescaled spectrum—these covariances can affect the statistical uncertainty by a factor of two or more

(see Fausnaugh 2017 for more details). We repeated this procedure  $10^3$  times and took the central 68% confidence interval of the output flux distributions as an estimate of the statistical uncertainty.

Because the integrated [O III]  $\lambda 5007$  line flux is not explicitly forced to be equal from night to night, the scatter of the [O III]  $\lambda 5007$  line light curve serves as an estimate of our calibration uncertainty (Barth et al. 2015). We extracted narrow [O III]  $\lambda 5007$  line light curves in the same way as for the broad lines, and the results are shown in Figure 7. Several points are noticeably below the means of their light curves, particularly for NGC 2617 and 3C 382. These observations were taken in poor weather and display significant scatter between the individual rescaled exposures prior to averaging. This suggests variable amounts of flux-losses between the AGN and extended [O III]  $\lambda 5007$ /host galaxy, due to variable seeing and large guiding errors that move the object in the slit. Although the rescaling model from Section 2.2.2 cannot correct this issue, the offsets of these points are not very large compared to the statistical uncertainties (no more than  $3.1\sigma$ ), and we opt to include them in the analysis. Since the effect due to spatially extended [O III]  $\lambda 5007$  emission is relatively small even in very poor conditions, it will be unimportant in good conditions.



**Figure 5.** Same as Figure 2 but for 3C 382. We do not use the H $\gamma$  or He II  $\lambda 4686$  light curves for the lag analysis in this study.

The fractional standard deviations of the narrow-line light curves are given in Table 1 and range between 0.1% and 1.4%. These values only represent our ability to correct for extrinsic variations (such as weather conditions) in the observed spectra. Additional systematic uncertainties dominate the epoch-to-epoch uncertainties of the light curves, including (but not limited to) the nightly sensitivity functions, continuum subtraction, and additional spectral components such as Fe II emission. The latter two issues are especially problematic for the He II  $\lambda 4686$  light curves.

To account for these systematics, we rescaled the light-curve uncertainties so that they approximate the observed flux variations from night to night. We selected three adjacent points,  $F(t_{j-1})$ ,  $F(t_j)$ , and  $F(t_{j+1})$ , linearly interpolated between  $F(t_{j-1})$  and  $F(t_{j+1})$ , and measure  $\Delta = [F(t_j) - I(t_j)]/\sigma(t_j)$ , where  $I(t_j)$  is the interpolated value at  $t_j$  and  $\sigma(t_j)$  is the statistical uncertainty on  $F(t_j)$ . The deviate  $\Delta$  therefore measures the departure of the light curve from a simple linear model. We calculated  $\Delta$  for  $j = 2$  to  $N_t - 1$  (i.e., ignoring the first and last points), and we multiplied the statistical uncertainties  $\sigma(t_j)$  by the mean absolute deviation (MAD)  $|\Delta|$ . We also imposed a minimum value of 1.0 on these rescaling factors. Inspection of the distribution of  $\Delta$  shows that the residuals are reasonably (but not perfectly) represented by a

Gaussian with a similar MAD value. This method ensures that the uncertainties account for any systematics that the rescaling model cannot capture. We have ignored the uncertainty in the interpolation  $I(t_j)$ , so our method slightly overestimates the required rescaling factors. Monte Carlo simulations may be able to assess the importance of uncertainty in  $I(t_j)$  for future work. The rescaling factors are given in Table 4 and are fairly small, generally running between 1.0 and 2.0, with a mean of 1.8 and a maximum of 3.42 for the H $\beta$  light curve in NGC 4051. NGC 4051 has the largest rescaling factors overall, which may be due to real short timescale variability that departs from our simple linear model (Denney et al. 2010). We therefore also experimented with using the unscaled light-curve uncertainties in our time-series analysis (Section 3) for this object. We found that our results do not sensitively depend on the scale of the uncertainties, although our Bayesian lag analysis (Section 3.2) indicates that the unscaled uncertainties are probably underestimated.

### 2.5.2. Broadband Light Curves

Differential photometric light curves were extracted from the subtracted broadband images using ISIS's built-in photometry package. The software performs PSF photometry by fitting a

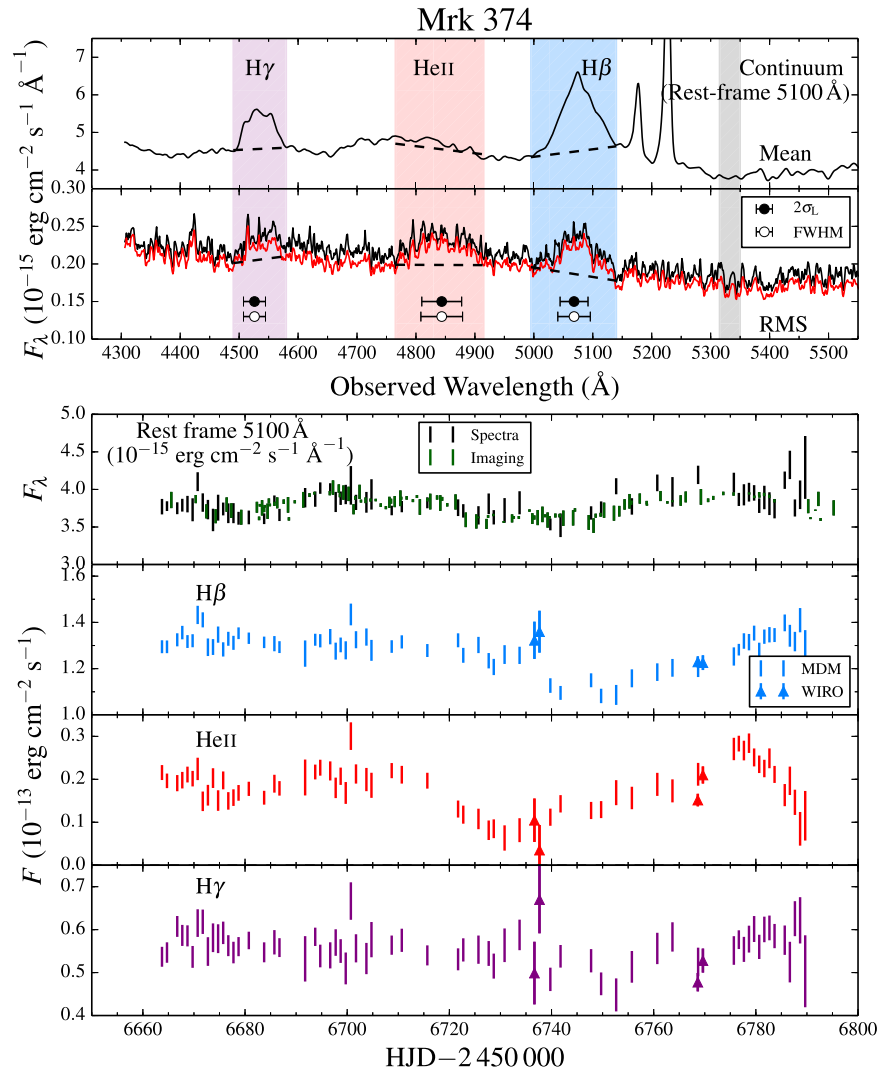
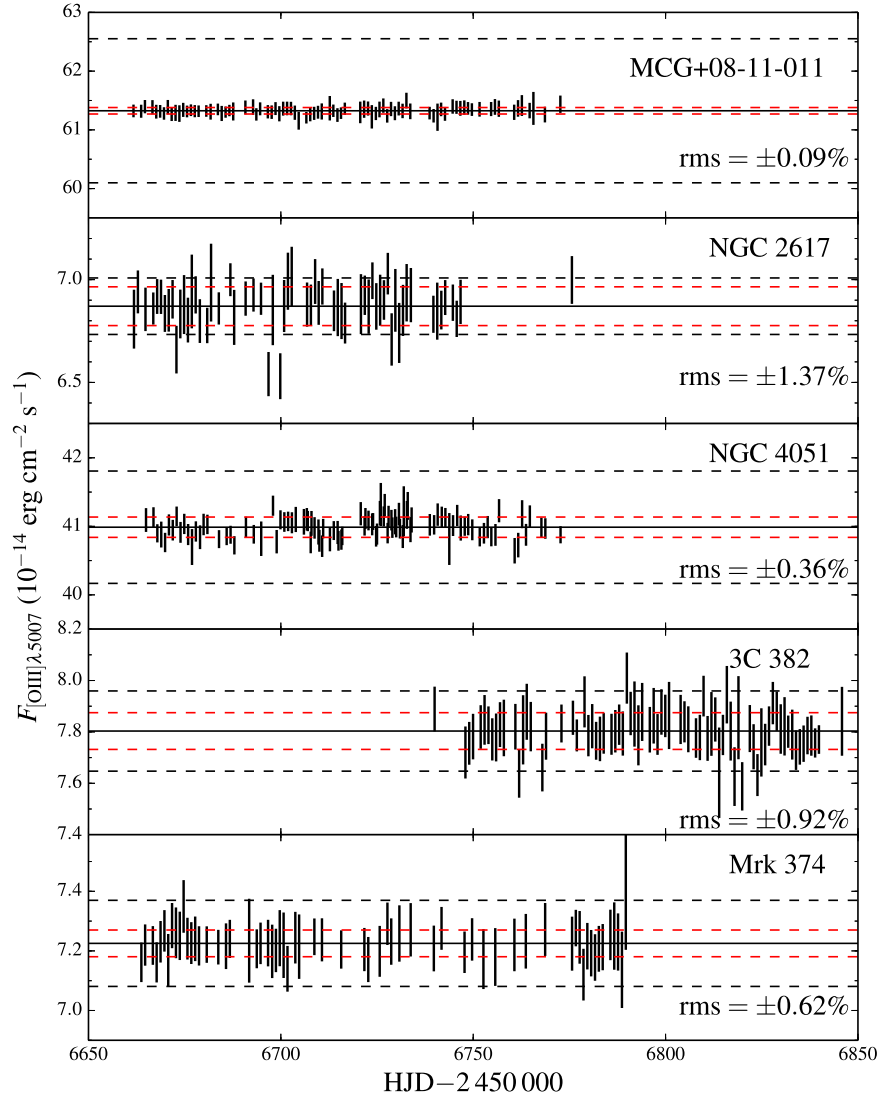


Figure 6. Same as Figure 2 but for Mrk 374.

model to the reference frame PSF and convolving this model with the kernel that was fitted during image subtraction. Because this transformation accounts for variable seeing, while the image subtraction has removed sources of constant flux, the output light curves cleanly isolate intrinsic variations of the AGN from contaminants such as host-galaxy starlight and seeing-dependent aperture effects. Any other constant systematic errors are also automatically subtracted out of the differential light curves. However, *ISIS* accounts for only the local Poisson uncertainty from photon-counting, while there are also systematic errors from imperfect subtractions (e.g., Hartman et al. 2004). We addressed this problem in the same way as Fausnaugh et al. (2016). We inspected the differential light curves of comparison stars and rescaled the uncertainties by a time dependent factor to make the comparison star residuals consistent with a constant model. The reduced  $\chi^2$  of the comparison star light curves is therefore set to one, which requires an average error rescaling factor of 1.0 to 5.0, depending on the object and the telescope. Since our targets are fairly bright, the formal *ISIS* uncertainties are very small and rescaling even by a factor of five results in uncertainties no greater than 3%–6%. See Section 2.2 of Fausnaugh et al. (2016) for more details.

We next calibrated the differential broadband light curves to the flux scale of the spectroscopic continuum light curve. The inter-calibration procedure solves for a maximum-likelihood shift and rescaling factor for each differential light curve, forcing the *V*-band photometry to match the rest-frame 5100 Å continuum flux. The inter-calibration parameters account for the different detector gains/bias levels, telescope throughputs, and (to first-order) a correction for the wider bandpass and different effective wavelengths of the broadband filters compared to the spectroscopic continuum averaging window. An advantage of this procedure is that it does not require accurate knowledge of the image zeropoints (or color corrections), which would otherwise limit the overall precision when combining data from different telescopes. The model also minimizes systematic errors that can result in strong correlations between measurements from the same telescope.

Because observations from various telescopes are never simultaneous, it is necessary to interpolate the light curves when fitting the inter-calibration parameters. We followed Fausnaugh et al. (2016) and modeled the time series as a damped random walk (DRW), as implemented by the *JAVELIN* software (Zu et al. 2011). Although recent studies have shown that the power spectra of AGN light curves on short timescales may be somewhat steeper than a DRW



**Figure 7.** Narrow [O III]  $\lambda 5007$  line light curves. The scatter of these light curves represents the precision of our final flux calibration. The dashed black lines show 2% variations around the mean, the dashed red lines show the rms scatter of the data, which range from 0.1% (MCG+08-11-011) to 1.4% (NGC 2617). Outliers are discussed in Section 2.5.1.

(Edelson et al. 2014; Kasliwal et al. 2015), Zu et al. (2013) found that the DRW is an adequate description of the timescales considered here (see also Skielboe et al. 2015; Fausnaugh et al. 2016; Kozłowski 2016a, 2016b). Our interpolation scheme and fitting procedure are identical to those described by Fausnaugh et al. (2016).

### 2.5.3. Light-curve Properties

The final light curves are shown in Figures 2–6 and given in Tables 5–14. We characterize the statistical properties of the light curves in Table 4, reporting the median cadence, mean flux level, and average S/N. We also measure the light-curve variability using a technique similar to our treatment of the variability spectra  $\sigma_{\text{var}}(\lambda)$ . In the presence of noise, it is necessary to separate the intrinsic variability from that due to measurement errors. We therefore define the intrinsic variability of the light curves as  $\sigma_{\text{var}}$  and solve for it by minimizing

$$-2 \ln \mathcal{L} = \sum_i \frac{[F(t_i) - \hat{F}]^2}{\sigma^2(t_i) + \sigma_{\text{var}}^2} + \sum_i \ln [\sigma^2(t_i) + \sigma_{\text{var}}^2], \quad (6)$$

where  $F(t_i)$  is the flux at epoch  $i$ ,  $\sigma(t_i)$  is its uncertainty, and  $\hat{F}$  is the optimal average flux (weighted by  $\sigma^2(t_i) + \sigma_{\text{var}}^2$ ). For small measurement uncertainties, the fractional variability  $\sigma_{\text{var}}/\hat{F}$  converges to the standard definition of the “excess variance” (Rodríguez-Pascual et al. 1997)

$$F_{\text{var}} = \frac{1}{\bar{F}} \sqrt{\frac{1}{N-1} \sum_i [F(t_i) - \bar{F}]^2 - \bar{\sigma}^2}, \quad (7)$$

where  $\bar{\sigma}$  is the time-averaged measurement uncertainty of the light curve. We therefore define  $F_{\text{var}} = \sigma_{\text{var}}/\hat{F}$  and report these values in Table 4. These values are slightly underestimated, since  $\hat{F}$  is not corrected for constant components (such as host-galaxy starlight or narrow-line emission). We also approximate the S/N of the variability as

$$(S/N)_{\text{var}} = \frac{\sigma_{\text{var}}}{\bar{\sigma} \sqrt{2/N_{\text{obs}}}}. \quad (8)$$

The  $\sqrt{2/N_{\text{obs}}}$  term enters because the variance of  $\bar{\sigma}$  is expected to approximately scale as that of a reduced  $\chi^2$  distribution.



**Table 4**  
Light-curve Properties

Object	Light curve	$N_{\text{obs}}$	$\Delta t_{\text{med}}$ (days)	Uncertainty Rescaling Factor	$\bar{F}$	$\langle S/N \rangle$	$F_{\text{var}}$	$(S/N)_{\text{var}}$	$r_{\text{max}}$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
MCG+08-11-011	5100 Å	190	0.59	1.51	4.49	71.2	0.10	67.4	$\equiv 1$
	H $\beta$	86	1.01	1.58	3.79	103.1	0.07	47.8	$0.90 \pm 0.01$
	H $\gamma$	82	1.01	1.52	1.29	34.2	0.09	19.2	$0.84 \pm 0.03$
	He II $\lambda$ 4686	86	1.01	1.46	0.31	6.8	0.44	19.6	$0.78 \pm 0.04$
NGC 2617	5100 Å	161	0.92	1.81	5.17	57.2	0.09	44.4	$\equiv 1$
	H $\beta$	61	1.01	1.91	3.31	39.5	0.10	21.1	$0.61 \pm 0.07$
	H $\gamma$	61	1.01	1.65	1.18	11.0	0.20	12.3	$0.62 \pm 0.07$
	He II $\lambda$ 4686	61	1.01	1.18	0.15	3.2	0.61	10.9	$0.49 \pm 0.08$
NGC 4051	5100 Å	270	0.47	1.00	12.90	191.7	0.02	49.7	$\equiv 1$
	H $\beta$	107	0.96	3.42	3.14	45.7	0.09	31.1	$0.59 \pm 0.05$
	H $\gamma$	98	0.99	2.48	1.92	26.1	0.08	13.7	$0.50 \pm 0.06$
	He II $\lambda$ 4686	107	0.96	3.25	1.59	12.9	0.11	10.2	$0.47 \pm 0.07$
3C 382	5100 Å	209	0.56	1.17	3.18	148.5	0.09	131.3	$\equiv 1$
	H $\beta$	81	1.00	1.70	2.06	43.4	0.05	14.5	$0.83 \pm 0.25$
	H $\gamma$	81	1.00	1.66	0.60	8.1	0.07	3.6	$0.59 \pm 0.24$
	He II $\lambda$ 4686	81	1.00	1.90	0.13	3.2	0.19	3.9	$0.61 \pm 0.47$
Mrk 374	5100 Å	180	0.59	2.39	3.80	94.5	0.03	25.5	$\equiv 1$
	H $\beta$	67	1.01	1.54	1.29	38.4	0.05	11.2	$0.50 \pm 0.07$
	H $\gamma$	67	1.01	1.79	0.56	18.7	0.04	3.8	$0.42 \pm 0.07$
	He II $\lambda$ 4686	67	1.01	1.41	0.18	8.2	0.25	11.9	$0.56 \pm 0.06$

**Note.** Column 3 gives the number of observations in each light curve. Column 4 gives the median cadence. Column 5 gives the rescaling factor by which the statistical uncertainties are multiplied to account for additional systematic errors (see Section 2.5.1). Column 6 gives the mean flux level of each light curve. The rest-frame 5100 Å continuum light curves are in units of  $10^{-15}$  erg cm $^{-2}$  s $^{-1}$  Å $^{-1}$ , and the emission-line light curves are in units of  $10^{-13}$  erg cm $^{-2}$  s $^{-1}$ . Column 7 gives the mean signal-to-noise ratio  $\langle S/N \rangle$ . Column 8 gives the rms fractional variability defined in Equation (6). Column 9 gives the approximate S/N at which we detect variability (see Section 2.5.3). Column 10 gives the maximum value of the interpolated cross-correlation function (see Section 3.1).

**Table 5**  
MCG+08-11-011 Continuum Light Curve

HJD (days)	$F_{\lambda}$ ( $10^{-15}$ ergs s $^{-1}$ cm $^{-2}$ Å $^{-1}$ )	Telescope ID
(1)	(2)	(3)
6639.5218	$3.6279 \pm 0.0448$	CrAO
6649.4750	$3.6413 \pm 0.0845$	CrAO
6653.4973	$3.8005 \pm 0.0631$	CrAO
6656.3942	$3.9176 \pm 0.0425$	CrAO
6661.6986	$3.9866 \pm 0.0764$	MDM
6662.2042	$3.9837 \pm 0.0412$	WC18

**Note.** Column 1 gives HJD—2,450,000 at mid-exposure. Column 2 give the continuum flux density and uncertainty. Column 3 identifies the contributing telescope (see Section 2.3).

(This table is available in its entirety in machine-readable form.)

However, this calculation assumes uncorrelated uncertainties, and a full analysis requires treatment of the red-noise properties of the light curve (see Vaughan et al. 2003).

With the exception of the 3C 382 H $\gamma$ , the 3C 382 He II  $\lambda$ 4686, and the Mrk 374 H $\gamma$  light curves, we detect variability in all of the other emission lines at greater than  $\sim 10\sigma$ . The variability amplitudes of MCG+08-11-011 and NGC 2617 are especially strong ( $F_{\text{var}} \gtrsim 10\%$ ). For NGC 4051, the continuum has little fractional variability ( $F_{\text{var}} = 2\%$ ), which may be caused by a high fraction of host-galaxy

**Table 6**  
NGC 2617 Continuum Light Curve

HJD (days)	$F_{\lambda}$ ( $10^{-15}$ ergs s $^{-1}$ cm $^{-2}$ Å $^{-1}$ )	Telescope ID
(1)	(2)	(3)
6639.6747	$4.6776 \pm 0.1089$	CrAO
6643.6320	$4.9717 \pm 0.1357$	CrAO
6644.5145	$5.0946 \pm 0.0871$	CrAO
6646.0981	$5.0904 \pm 0.0892$	WC18
6646.5287	$5.4647 \pm 0.0885$	CrAO
6647.0990	$5.3930 \pm 0.1254$	WC18

**Note.** Columns are the same as in Table 5.

(This table is available in its entirety in machine-readable form.)

**Table 7**  
NGC 4051 Continuum Light Curve

HJD (days)	$F_{\lambda}$ ( $10^{-15}$ ergs s $^{-1}$ cm $^{-2}$ Å $^{-1}$ )	Telescope ID
(1)	(2)	(3)
6645.6113	$12.7318 \pm 0.0624$	WC18
6646.6098	$12.8314 \pm 0.0668$	WC18
6647.6275	$13.0803 \pm 0.0683$	WC18
6648.5919	$12.9598 \pm 0.0499$	WC18
6650.5178	$12.8375 \pm 0.0550$	WC18
6653.5959	$13.0814 \pm 0.0956$	WC18

**Note.** Columns are the same as in Table 5.

(This table is available in its entirety in machine-readable form.)

**Table 8**  
3C 382 Continuum Light Curve

HJD (days) (1)	$F_{\lambda}$ ( $10^{-15}$ ergs s $^{-1}$ cm $^{-2}$ Å $^{-1}$ ) (2)	Telescope ID (3)
6670.6411	$3.2490 \pm 0.1390$	WC18
6689.0237	$3.0685 \pm 0.0094$	LCOGT1
6690.0279	$3.0018 \pm 0.0105$	LCOGT1
6691.6169	$3.0619 \pm 0.1197$	WC18
6693.0025	$2.9948 \pm 0.0106$	LCOGT1
6696.9897	$2.8652 \pm 0.0094$	LCOGT1

**Note.** Columns are the same as in Table 5.

(This table is available in its entirety in machine-readable form.)

**Table 9**  
Mrk 374 Continuum Light Curve

HJD (days) (1)	$F_{\lambda}$ ( $10^{-15}$ ergs s $^{-1}$ cm $^{-2}$ Å $^{-1}$ ) (2)	Telescope ID (3)
6663.7432	$3.7706 \pm 0.1030$	MDM
6664.7221	$3.7776 \pm 0.0649$	MDM
6665.5588	$3.8676 \pm 0.1114$	WC18
6666.7292	$3.7496 \pm 0.0791$	MDM
6667.7164	$3.7902 \pm 0.0850$	MDM
6668.7316	$3.8045 \pm 0.1097$	MDM

**Note.** Columns are the same as in Table 5.

(This table is available in its entirety in machine-readable form.)

starlight. For MCG+08-11-011, NGC 2617, and NGC 4051, the median cadence is near 1 day for all light curves, and the mean S/N usually ranges from several tens to hundreds. In fact, the S/N in the spectra is even higher, reaching 100 to 300 per pixel in the continuum. Combined with the large variability amplitudes, it is likely that we will be able to construct velocity-delay maps and dynamical models for these objects in future work.

### 3. Time-series Measurements

We measure lags between continuum and line light curves using two independent methods: traditional cross-correlation techniques and a Bayesian analysis using the JAVELIN software.

#### 3.1. Cross Correlation

The cross-correlation procedure derives a lag from the centroid of the interpolated cross-correlation function (ICCF, Gaskell & Peterson 1987), as implemented by Peterson et al. (2004). For a given time delay, we shift the abscissas of the first light curve, linearly interpolate the second light curve to the new time coordinates, and calculate the correlation coefficient  $r_{cc}$  between all overlapping data points. We then repeat this calculation but shift the second light curve by the negative of the given time delay and interpolate the first light curve. The two values of  $r_{cc}$  are averaged together and the ICCF is evaluated by repeating this procedure on a grid of time delays spaced by 0.1 days. All ICCFs are measured relative to the 5100 Å continuum light curve (intercalibrated with the broadband measurements). For each line light curve, the maximum value  $r_{\max}$  of the ICCF is given in Table 4.

The lag is estimated with the ICCF centroid, defined as  $\tau_{\text{cent}} = \int \tau r_{cc}(\tau) d\tau / \int r_{cc}(\tau) d\tau$  for values of  $r_{cc} \geq 0.8r_{\max}$ .

We estimate the uncertainty on  $\tau_{\text{cent}}$  using the flux randomization/random subset sampling (FR/RSS) method of Peterson et al. (2004). This technique generates perturbed light curves by randomly selecting (with replacement) a subset of the data from both light curves and adjusting the fluxes by a Gaussian deviate scaled to the measurement uncertainties. The lag  $\tau_{\text{cent}}$  is calculated for  $10^3$  perturbations of the data and its uncertainty is estimated from the central 68% confidence interval of the resulting distribution. The ICCF and centroid distributions are shown in Figure 8 for all objects and line light curves, and Table 15 gives the median values and central 68% confidence intervals of these distributions. For completeness, we also report in Table 15 the lag  $\tau_{\text{peak}}$  that corresponds to  $r_{\max}$ . Note that these lags have been corrected to the rest frame of the source. For 3C 382, we do not find meaningful centroids in the ICCFs of the H $\gamma$  and He II  $\lambda 4686$  light curves. This is because of the width of the autocorrelation function of the continuum and its poor correlation with the line light curves. We therefore do not include these lines for the rest of the ICCF analysis.

Long-term trends in the light curves can bias the resulting ICCF due to red-noise leakage (Welsh 1999). We therefore experimented with detrending the light curves and/or restricting the baseline over which to calculate the ICCF. For MCG +08-11-011, these experiments had no effect, while for Mrk 374 and 3C 382 they eliminated any lag signal in the data. For NGC 2617, we found that restricting the data to  $6620 < \text{HJD} - 2,450,000 < 6730$  improved the ICCF by narrowing the central peak, as shown in the top four panels of Figure 9. However, this restriction changed the ICCF centroid by only 0.01 days, a negligible amount. For NGC 2617, the peaks in the H $\gamma$  and He II  $\lambda 4686$  ICCFs at  $\pm 25$  days are also obvious aliases, so we only report the lag based on the peak near 0 days. For NGC 4051, we found that detrending the continuum and line light curves with a second-order polynomial improves the ICCF, as shown in the bottom four panels of Figure 9. The long-term continuum trend is very weak, but there is a strong positive trend in the line light curves that is dominated by the linear term. Subtracting this linear trend decreases the median of the centroid distribution from 4.92 to 2.56 days, a change of  $1.5\sigma$ . We adopt the smaller lag because of the quality of the detrended ICCF, and our Bayesian method (described below) finds a lag consistent with this smaller value.

#### 3.2. JAVELIN

We also investigated the line lags using a Bayesian approach, as implemented by the JAVELIN software (Zu et al. 2011). JAVELIN explicitly models the reverberating light curves and corresponding transfer functions so as to find a posterior probability distribution of lags. We have already discussed JAVELIN's assumption that light curves are reasonably characterized by a DRW (Section 2.5.2). JAVELIN also assumes that the transfer function is a simple top-hat that can be parameterized by a width, an amplitude, and a mean time delay. This assumption is not very restrictive, since it is difficult to distinguish among transfer functions in the presence of noise (Rybicki & Kleyna 1994; Zu et al. 2011) and a top-hat is broadly consistent with expectations for physically plausible BLR geometries (e.g., disks or spherical shells).

**Table 10**  
MCG+08-11-011 Line Light Curves

HJD	H $\beta$	H $\gamma$	He II	Telescope ID
(days)		( $10^{-13}$ erg s $^{-1}$ cm $^{-2}$ )		
(1)	(2)	(3)	(4)	(5)
6661.6986	3.4660 $\pm$ 0.0288	1.2156 $\pm$ 0.0298	0.1923 $\pm$ 0.0416	MDM
6663.6924	3.4891 $\pm$ 0.0302	1.1841 $\pm$ 0.0291	0.2341 $\pm$ 0.0401	MDM
6664.6734	3.4978 $\pm$ 0.0346	1.1986 $\pm$ 0.0326	0.2352 $\pm$ 0.0407	MDM
6666.6277	3.4723 $\pm$ 0.0387	1.0566 $\pm$ 0.0388	0.3973 $\pm$ 0.0505	MDM
6667.6147	3.5429 $\pm$ 0.0365	1.1637 $\pm$ 0.0354	0.4047 $\pm$ 0.0471	MDM
6668.6285	3.5031 $\pm$ 0.0307	1.2122 $\pm$ 0.0333	0.3580 $\pm$ 0.0437	MDM

**Note.** Column 1 gives HJD–2,450,000 at mid-exposure. Columns 2–4 give the line fluxes and their uncertainties. Column 5 identifies the contributing telescope (see Section 2.3).

(This table is available in its entirety in machine-readable form.)

**Table 11**  
NGC 2617 Line Light Curves

HJD	H $\beta$	H $\gamma$	He II	Telescope ID
(days)		( $10^{-13}$ erg s $^{-1}$ cm $^{-2}$ )		
(1)	(2)	(3)	(4)	(5)
6661.8270	3.4179 $\pm$ 0.1151	1.2480 $\pm$ 0.1362	0.0439 $\pm$ 0.0614	MDM
6662.8456	3.4662 $\pm$ 0.0848	1.2049 $\pm$ 0.1039	0.0145 $\pm$ 0.0470	MDM
6664.8143	3.2528 $\pm$ 0.0756	1.0938 $\pm$ 0.0892	0.0779 $\pm$ 0.0370	MDM
6666.8417	3.3192 $\pm$ 0.0635	1.1110 $\pm$ 0.0744	0.1912 $\pm$ 0.0346	MDM
6667.8446	3.3112 $\pm$ 0.0621	1.2235 $\pm$ 0.0793	0.1744 $\pm$ 0.0363	MDM
6668.8507	3.2390 $\pm$ 0.0690	1.3410 $\pm$ 0.0837	0.1203 $\pm$ 0.0344	MDM

**Note.** Columns are the same as in Table 10.

(This table is available in its entirety in machine-readable form.)

**Table 12**  
NGC 4051 Line Light Curves

HJD	H $\beta$	H $\gamma$	He II	Telescope ID
(days)		( $10^{-13}$ erg s $^{-1}$ cm $^{-2}$ )		
(1)	(2)	(3)	(4)	(5)
6664.9552	2.9038 $\pm$ 0.0705	1.8998 $\pm$ 0.0729	1.3355 $\pm$ 0.1207	MDM
6666.8941	2.7895 $\pm$ 0.0616	1.8579 $\pm$ 0.0659	1.4384 $\pm$ 0.1032	MDM
6667.8962	2.7808 $\pm$ 0.0565	1.7772 $\pm$ 0.0605	1.3202 $\pm$ 0.0981	MDM
6668.9146	2.8279 $\pm$ 0.0708	1.8106 $\pm$ 0.0767	1.6972 $\pm$ 0.1251	MDM
6669.9111	2.8152 $\pm$ 0.0586	1.7299 $\pm$ 0.0599	1.5860 $\pm$ 0.1034	MDM
6670.9090	2.8217 $\pm$ 0.0657	1.8284 $\pm$ 0.0694	1.5306 $\pm$ 0.1166	MDM

**Note.** Columns are the same as in Table 10.

(This table is available in its entirety in machine-readable form.)

**Table 13**  
3C 382 Line Light Curves

HJD	H $\beta$	H $\gamma$	He II	Telescope ID
(days)		( $10^{-13}$ erg s $^{-1}$ cm $^{-2}$ )		
(1)	(2)	(3)	(4)	(5)
6739.9650	2.0891 $\pm$ 0.0519	0.6969 $\pm$ 0.0820	0.1308 $\pm$ 0.0429	MDM
6747.9747	2.0787 $\pm$ 0.0581	0.6508 $\pm$ 0.0864	0.1038 $\pm$ 0.0452	MDM
6748.9640	2.0885 $\pm$ 0.0447	0.6067 $\pm$ 0.0715	0.0901 $\pm$ 0.0388	MDM
6749.9571	2.0675 $\pm$ 0.0487	0.6094 $\pm$ 0.0785	0.0745 $\pm$ 0.0392	MDM
6751.9471	2.0665 $\pm$ 0.0520	0.6116 $\pm$ 0.0784	0.1327 $\pm$ 0.0392	MDM
6752.9466	2.0513 $\pm$ 0.0550	0.4939 $\pm$ 0.0833	0.1497 $\pm$ 0.0464	MDM

**Note.** Columns are the same as in Table 10.

(This table is available in its entirety in machine-readable form.)

**Table 14**  
Mrk 374 Line Light Curves

HJD	H $\beta$	H $\gamma$	He II	Telescope ID
(days)		( $10^{-13}$ erg s $^{-1}$ cm $^{-2}$ )		
(1)	(2)	(3)	(4)	(5)
6663.7432	$1.2948 \pm 0.0280$	$0.5368 \pm 0.0232$	$0.2157 \pm 0.0174$	MDM
6664.7221	$1.2944 \pm 0.0274$	$0.5467 \pm 0.0238$	$0.1965 \pm 0.0170$	MDM
6666.7292	$1.3247 \pm 0.0282$	$0.6074 \pm 0.0243$	$0.1911 \pm 0.0188$	MDM
6667.7164	$1.3559 \pm 0.0291$	$0.5865 \pm 0.0245$	$0.1980 \pm 0.0193$	MDM
6668.7316	$1.3172 \pm 0.0278$	$0.5855 \pm 0.0245$	$0.2111 \pm 0.0184$	MDM
6669.7373	$1.3207 \pm 0.0282$	$0.5365 \pm 0.0254$	$0.2015 \pm 0.0181$	MDM

**Note.** Columns are the same as in Table 5.

(This table is available in its entirety in machine-readable form.)

We ran JAVELIN models for each line using the 5100 Å continuum as the driving light curve, and we used internal JAVELIN routines to remove any linear trends from the light curves during the fit. The damping timescale (a parameter of the DRW model) for most AGNs is several hundred days or longer (Kelly et al. 2009; MacLeod et al. 2010), and our light curves are not long enough to meaningfully constrain this parameter. We therefore (arbitrarily) fixed the damping timescale to 200 days. We also tested several different damping timescales (from a few days to 500 days), and found that the choice of 200 days does not affect the best-fit lags—an exact estimate of the damping timescale is not necessary to reasonably interpolate the light curves (Kozłowski 2016b). Table 15 gives the median and 68% confidence interval of the posterior lag distributions, denoted as  $\tau_{\text{JAV}}$ . We also employed models that fit all light curves from a single object simultaneously, which maximizes the available information. These results are given in Table 15 as  $\tau_{\text{multi}}$ . Posterior distributions of  $\tau_{\text{multi}}$  are shown by the blue histograms in Figure 8. For the H $\gamma$  and He II  $\lambda 4686$  light curves from 3C 382, we were again unable to constrain any lag signal, and we drop these light curves from the rest of this analysis.

### 3.3. Results

We generally find consistent results between the ICCF method and JAVELIN models. The largest discrepancies are the H $\beta$  lags for NGC 2617 ( $\Delta\tau = 1.6\sigma$ ) and 3C 382 ( $\Delta\tau = 2.0\sigma$ ), but these differences are not statistically significant. In NGC 2617, where the ICCF method detects a lag consistent with zero in the H $\gamma$  or He II  $\lambda 4686$  light curves, JAVELIN finds a lag at reasonably high confidence: the percentiles for  $\tau_{\text{multi}} = 0$  in the posterior lag distributions of H $\gamma$  and He II  $\lambda 4686$  are 8.3% and 1.1%, which are  $1.4\sigma$  and  $2.3\sigma$  detections for Gaussian probability distributions, respectively. For Mrk 374, a H $\gamma$  lag is detected at high significance using JAVELIN (we do not claim a lag detection for He II  $\lambda 4686$  in this object, since the  $\tau_{\text{multi}} = 0$  percentile is 20%, only  $0.2\sigma$  for a Gaussian probability distribution). The detection of these lags represents a significant advantage of the JAVELIN technique over traditional cross-correlation methods. We adopt the  $\tau_{\text{multi}}$  as our final lag measurements, since the multi-line global fits provide well-constrained lags, properly treat covariances between the lags from different light curves, and utilize the maximum amount of information available in the data.

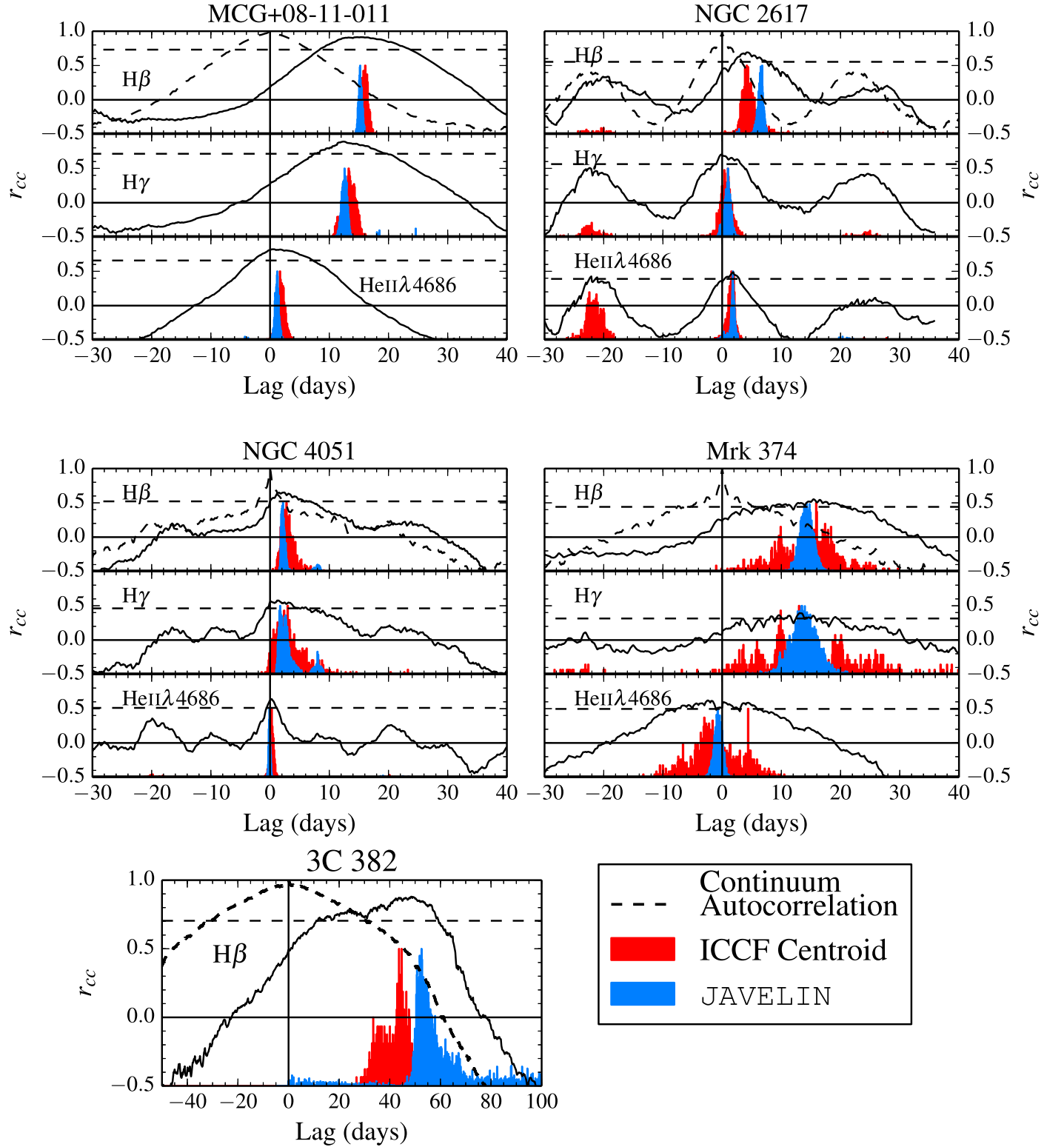
The analysis of NGC 4051 is especially difficult because the light curves exhibit low-amplitude variations. The lags in this object are also expected to be small, based on the AGN luminosity (Bentz et al. 2013) and a previous well-sampled RM experiment (Denney et al. 2009b). For H $\beta$ , JAVELIN finds a definite lag near 2 days, consistent with the detrended ICCF approach. For H $\gamma$ , the ICCF method finds a lag consistent with zero, while the single-line JAVELIN fit finds a lag of  $4.87 \pm 0.18$  days and the multi-line fit finds a lag of  $2.40 \pm 0.80$  days (rest frame). The single-line fit results in a complicated multi-modal posterior distribution with smaller peaks at 15 and 25 days that are caused by aliasing. For example, the 25 day lag is probably caused by aligning the H $\gamma$  maximum near 6745 days with the local maximum in the continuum light curve at 6720 days (Figure 4). However, the multi-line fit shows a strong, dominant peak for H $\gamma$  at 2.40 days (rest frame). A probable explanation is that the H $\beta$  light curve matches the overall shape of H $\gamma$ , but has stronger features against which to estimate a continuum lag—fitting both light curves simultaneously can therefore establish an H $\gamma$  lag with higher confidence. The problem with the H $\gamma$  light curve appears in a more serious form in the He II  $\lambda 4686$  light curve, and JAVELIN finds a lag consistent with zero for this line.

### 4. Linewidths and $M_{\text{BH}}$ Calculations

After determining the characteristic size of the BLR from the mean time delay, the next step is to calculate the characteristic line-of-sight velocity of the BLR gas, from which we can derive SMBH masses. The BLR velocity is estimated from the width of emission lines in the MDM spectra. However, it is important to use the linewidth of the variable component of the profile, since we measure the BLR radius from the variable line flux. For example, the variable profile of 3C 382 is radically different (and much broader) than the time-averaged profile in the mean spectrum (Figure 5). We therefore measure and report in Table 16 linewidths both in the mean spectrum  $\hat{F}(\lambda)$ , and in the rms spectrum  $\sigma_{\text{var}}(\lambda)$ , but we use the latter for mass determinations.

There are two common choices for linewidth measurements: the full-width at half maximum (FWHM) and the line dispersion  $\sigma_L$  (the rms width of the line profile). There are advantages and disadvantages associated with both approaches—while the FWHM is simpler to measure, there are ambiguities for noisy or complicated line profiles such as the double-peaked H $\beta$





**Figure 8.** Lags for all emission lines in all objects. The solid lines show the ICCFs and the dashed lines show the autocorrelation function of the continuum. The red histograms show the ICCF centroid distributions for  $\tau_{cent}$  and the blue histograms show the posterior JAVELIN lag distributions for  $\tau_{multi}$ . The histograms are normalized by dividing by their maximum values. The horizontal dashed lines show  $0.8r_{max}$ , used to calculate  $\tau_{cent}$ .

profiles in MCG+08-11-011, NGC 2617, and 3C 382. On the other hand, although  $\sigma_L$  is well defined for arbitrary line profiles, it depends more sensitively on continuum subtraction and blending in the line wings (Denney et al. 2016; Mejía-Restrepo

et al. 2016). Peterson et al. (2004) found that velocities estimated with  $\sigma_L$  produce a tighter virial relation, and Denney et al. (2013) found that the masses determined from UV and optical lines agree better using  $\sigma_L$ . We therefore adopt  $\sigma_L$  as a measure of the

**Table 15**  
Rest-frame Line Lags

Object	Line	$\tau_{\text{cent}}$ (days)	$\tau_{\text{peak}}$ (days)	$\tau_{\text{JAV}}$ (days)	$\tau_{\text{multi}}$ (days)
(1)	(2)	(3)	(4)	(5)	(6)
MCG+08-11-011	H $\beta$	$15.72^{+0.50}_{-0.32}$	$15.02^{+1.86}_{-1.08}$	$15.06^{+0.26}_{-0.28}$	$14.98^{+0.34}_{-0.28}$
	H $\gamma$	$13.14^{+1.12}_{-1.05}$	$12.08^{+0.69}_{-0.98}$	$11.92^{+0.44}_{-0.44}$	$12.38^{+0.46}_{-0.49}$
	He II $\lambda 4686$	$1.88^{+0.58}_{-0.64}$	$1.59^{+1.27}_{-0.98}$	$1.24^{+0.36}_{-0.29}$	$1.21^{+0.29}_{-0.33}$
NGC 2617	H $\beta$	$4.32^{+1.10}_{-1.35}$	$4.16^{+1.08}_{-1.68}$	$6.22^{+0.51}_{-0.54}$	$6.38^{+0.44}_{-0.50}$
	H $\gamma$	$0.91^{+1.50}_{-1.08}$	$0.61^{+1.28}_{-0.89}$	$0.83^{+0.58}_{-0.59}$	$0.81^{+0.59}_{-0.61}$
	He II $\lambda 4686$	$1.59^{+0.49}_{-0.69}$	$1.79^{+0.20}_{-0.89}$	$1.78^{+0.30}_{-0.35}$	$1.75^{+0.34}_{-0.38}$
NGC 4051	H $\beta$	$2.87^{+0.86}_{-1.33}$	$2.42^{+1.00}_{-1.90}$	$2.41^{+0.38}_{-0.46}$	$2.24^{+0.39}_{-0.28}$
	H $\gamma$	$2.82^{+1.82}_{-2.84}$	$2.82^{+1.90}_{-3.09}$	$4.87^{+0.28}_{-0.08}$	$2.40^{+0.86}_{-0.73}$
	He II $\lambda 4686$	$0.27^{+0.33}_{-0.40}$	$0.23^{+0.40}_{-0.40}$	$0.06^{+0.17}_{-0.17}$	$-0.03^{+0.16}_{-0.15}$
3C 382	H $\beta$	$40.49^{+8.02}_{-3.74}$	$43.58^{+4.16}_{-3.50}$	$52.07^{+3.18}_{-9.46}$	...
Mrk 374	H $\beta$	$14.84^{+5.76}_{-3.30}$	$14.81^{+5.85}_{-3.55}$	$15.03^{+1.41}_{-1.26}$	$13.73^{+1.06}_{-1.02}$
	H $\gamma$	$12.31^{+9.82}_{-9.80}$	$12.51^{+9.69}_{-11.80}$	$15.44^{+3.26}_{-2.85}$	$13.37^{+2.11}_{-2.08}$
	He II $\lambda 4686$	$-1.53^{+3.21}_{-5.79}$	$-1.59^{+2.69}_{-5.85}$	$-0.44^{+0.71}_{-0.68}$	$-0.57^{+0.65}_{-0.64}$

**Note.** Columns 3 and 4 give the centroids and peaks, respectively, of the interpolated cross-correlation functions (ICCFs). The uncertainties give the central 68% confidence intervals of the ICCF distributions from the FR/RSS procedure (see Section 3.1). Column 5 gives the lag fit by JAVELIN. Column 6 gives the same but using all light curves from a single object simultaneously. The uncertainties give the central 68% confidence intervals of the JAVELIN posterior lag distributions. All lags are relative to the 5100 Å continuum light curve and corrected to the rest frame. The uncertainties only represent the statistical errors—choices of continuum windows, detrending procedures, etc., introduce additional systematic uncertainties.

BLR velocity in this study. For completeness, we also give the FWHM in Table 16.

Linewidth uncertainties are estimated using a bootstrapping method. For  $10^3$  iterations on each object with  $N$  nightly spectra, we randomly select  $N$  observations with replacement, recompute the mean and rms spectrum, and remeasure the linewidths in the rms spectrum. The central 68% confidence interval of the resulting distributions are adopted as the formal uncertainty of the linewidth. This approach can only account for statistical uncertainties in the linewidths, which therefore represent lower limits on the uncertainties. There are additional systematic errors from the choice of wavelength windows that define the line profiles (Tables 2 and 3), as well as blending of the broad-line wings. The choice of wavelength windows and continuum subtraction is problematic for weak lines, lines with low variability, and lines with unusual profiles. In particular, our estimates for the He II  $\lambda 4686$  line in NGC 2617, NGC 4051, 3C 382, and all lines in Mrk 374 are certainly affected. Furthermore, the blue wing of H $\beta$  and the red wing of He II  $\lambda 4686$  overlap in MCG+08-11-011 and NGC 2617, and it is likely that the He II  $\lambda 4686$  velocity is severely underestimated (the effect on H $\beta$  is probably smaller, though it may not be negligible). Spectral decompositions may help with these problems in future analyses; for now, we note that the linewidth uncertainties are underestimated in these cases and we provide a treatment for this issue below.

We correct the linewidth measurements for the instrument resolution by subtracting the rms width of the spectrograph’s line-spread-function (LSF) in quadrature from the observed value of  $\sigma_L$ . Previous studies have found that the width of the LSF for the MDM spectrograph is near 3.2 or 3.4 Å (FWHM 7.6–7.9 Å, Denney et al. 2010; Grier et al. 2012). Based on comparisons with high spectral resolution observations, where the LSF width is negligible, we find an LSF width of 2.97 Å (FWHM = 6.99 Å). This value was determined using the

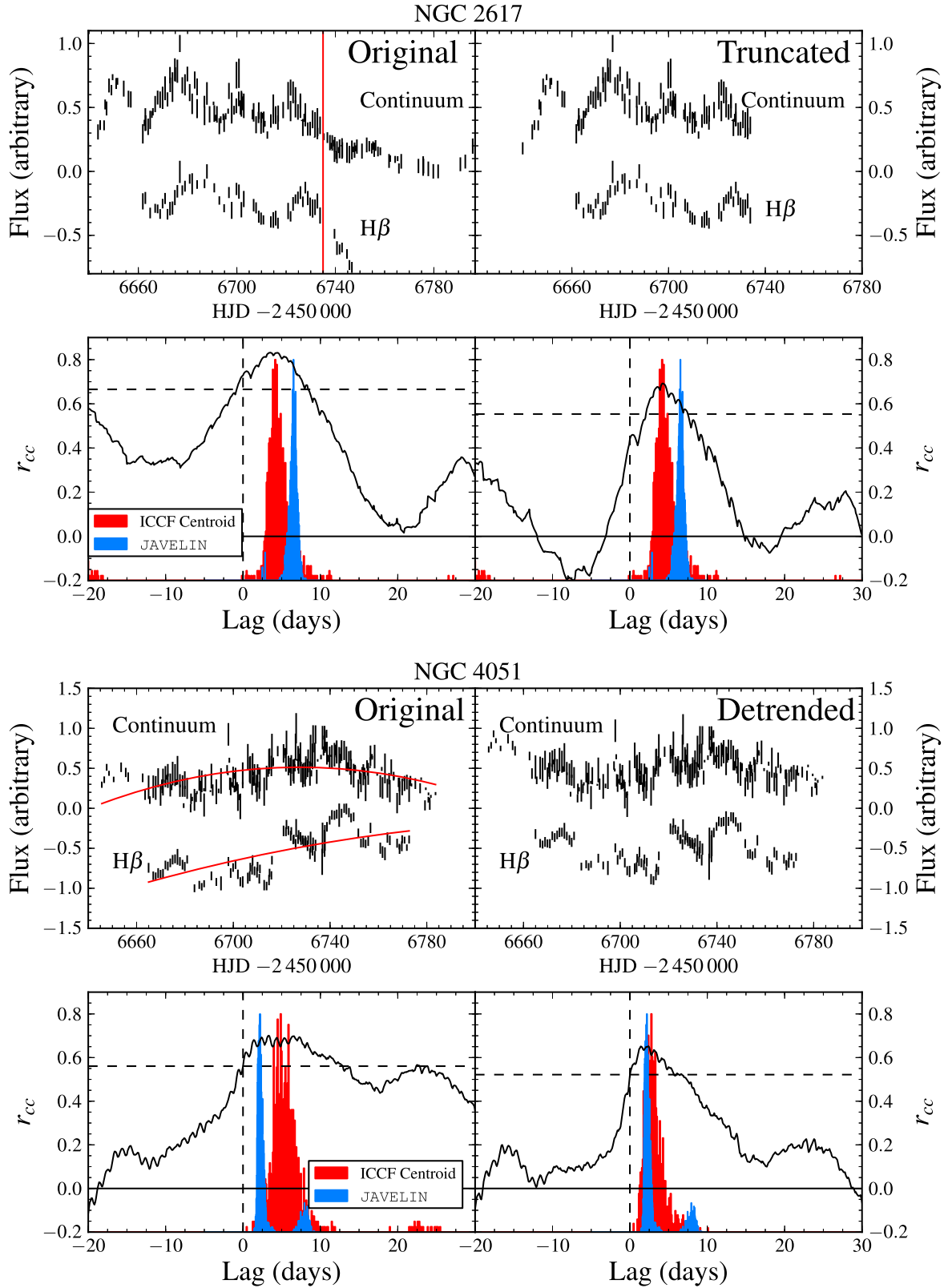
catalog of high-resolution [O III]  $\lambda 5007$  measurements from Whittle (1992), which contains intrinsic [O III]  $\lambda 5007$  line-widths for MCG+0-11-011 and NGC 4051. The [O III]  $\lambda 5007$  line of NGC 4051 is undersampled in the MDM spectra (the intrinsic FWHM is 190 km s $^{-1}$ , or 3.16 Å in the observed frame), and does not give a reliable estimate the instrumental broadening. However, the intrinsic [O III]  $\lambda 5007$  FWHM in MCG+08-11-011 is 605 km s $^{-1}$ , or 10.52 Å in the observed frame, which is well resolved. The observed FWHM in the MCG+08-11-011 reference spectrum (before smoothing, see Section 2.2.2 and below) is 12.63 Å, which implies that the FWHM of the LSF is 6.99 Å (a rms width of 2.97 Å). This value is close to but slightly smaller than previous estimates. The MDM LSF may not be perfectly stable in time, so we adopt 2.97 Å as the rms width of the instrumental broadening in our observations.

An additional correction must be applied because we smooth our reference spectra to approximately match the nights with the worst spectroscopic resolution (see Section 2.2.2). The kernel widths for this smoothing procedure were 1.4 Å for MCG+08-11-011, 1.5 Å for NGC 2617, 1.8 Å for NGC 4051, 1.7 Å for 3C 382, and 1.9 Å for Mrk 374 (the FWHM values are a factor of 2.35 larger). We also subtract these values in quadrature from the observed line dispersion. The final rest-frame linewidths and their uncertainties are given in Table 16.

We measure the SMBH masses as

$$M_{\text{BH}} = \langle f \rangle \frac{\sigma_L^2 c \tau_{\text{multi}}}{G}, \quad (9)$$

where  $c$  is the speed of light,  $G$  is the gravitational constant, and  $\langle f \rangle$  is the virial factor. The virial factor accounts for the unknown geometry and dynamics of the BLR and is determined by calibrating a sample of RM AGNs to the  $M_{\text{BH}}\text{--}\sigma_*$  relation (e.g., Onken et al. 2004; Park et al. 2012a;



**Figure 9.** Comparison of two CCF analyses of the H $\beta$  light curves for NGC 2617 and NGC 4051. NGC 2617 is shown in the top four panels and NGC 4051 is shown in the bottom four. Color coding of the histograms is the same as Figure 8. The left columns show the original light curves, the right columns show the modified light curves. In the case of NGC 2617, truncating the light curves at HJD-2,450,000 = 6 735 days (vertical red line, top left) helps concentrate the ICCF peak. For NGC 4051, subtracting a second-order polynomial fit to the data removes the effect of long-term secular trends (red lines, bottom-middle left panel). As a reminder, the JAVELIN fits include linear detrending, so the blue histograms are identical on the left and right sides. The same procedures were applied to all emission-line light curves in these objects.

**Table 16**  
Rest-frame Velocity Linewidth Measurements

Object	Line	rms Spectrum		Mean Spectrum	
		$\sigma_L$ (km s <sup>-1</sup> )	FWHM (km s <sup>-1</sup> )	$\sigma_L$ (km s <sup>-1</sup> )	FWHM (km s <sup>-1</sup> )
(1)	(2)	(3)	(4)	(5)	(6)
MCG +08- 11- 011	H $\beta$	1466 <sup>+102</sup> <sub>-174</sub>	4475 <sup>+192</sup> <sub>-356</sub>	1681 <sup>+02</sup> <sub>-02</sub>	1159 <sup>+0008</sup> <sub>-0007</sub>
	H $\gamma$	1604 <sup>+083</sup> <sub>-082</sub>	3916 <sup>+616</sup> <sub>-716</sub>	1175 <sup>+07</sup> <sub>-07</sub>	1978 <sup>+0026</sup> <sub>-0019</sub>
	He II $\lambda$ 4686	2453 <sup>+125</sup> <sub>-130</sub>	6617 <sup>+993</sup> <sub>-776</sub>	2893 <sup>+42</sup> <sub>-37</sub>	2517 <sup>+1814</sup> <sub>-1020</sub>
NGC 2617	H $\beta$	2424 <sup>+091</sup> <sub>-086</sub>	6489 <sup>+213</sup> <sub>-162</sub>	2709 <sup>+006</sup> <sub>-006</sub>	5303 <sup>+49</sup> <sub>-46</sub>
	H $\gamma$	3084 <sup>+086</sup> <sub>-090</sub>	7674 <sup>+423</sup> <sub>-472</sub>	2385 <sup>+033</sup> <sub>-034</sub>	4101 <sup>+26</sup> <sub>-29</sub>
	He II $\lambda$ 4686	2020 <sup>+329</sup> <sub>-572</sub>	6788 <sup>+984</sup> <sub>-855</sub>	3113 <sup>+147</sup> <sub>-218</sub>	7150 <sup>+35</sup> <sub>-33</sub>
NGC 4051	H $\beta$	493 <sup>+34</sup> <sub>-36</sub>	941 <sup>+017</sup> <sub>-019</sub>	470 <sup>+2</sup> <sub>-2</sub>	765 <sup>+03</sup> <sub>-03</sub>
	H $\gamma$	641 <sup>+55</sup> <sub>-59</sub>	1098 <sup>+031</sup> <sub>-034</sub>	942 <sup>+4</sup> <sub>-4</sub>	1676 <sup>+06</sup> <sub>-05</sub>
	He II $\lambda$ 4686	1689 <sup>+36</sup> <sub>-38</sub>	3885 <sup>+299</sup> <sub>-218</sub>	1898 <sup>+4</sup> <sub>-3</sub>	4598 <sup>+10</sup> <sub>-11</sub>
3C 382	H $\beta$	4552 <sup>+0214</sup> <sub>-0163</sub>	11549 <sup>+1292</sup> <sub>-0667</sub>	3227 <sup>+07</sup> <sub>-07</sub>	3619 <sup>+282</sup> <sub>-050</sub>
	H $\gamma$	5083 <sup>+1114</sup> <sub>-1942</sub>	10706 <sup>+1050</sup> <sub>-1294</sub>	2845 <sup>+55</sup> <sub>-53</sub>	3483 <sup>+022</sup> <sub>-022</sub>
	He II $\lambda$ 4686	2073 <sup>+0170</sup> <sub>-0352</sub>	4374 <sup>+1257</sup> <sub>-0602</sub>	1789 <sup>+27</sup> <sub>-28</sub>	5186 <sup>+247</sup> <sub>-075</sub>
Mrk 374	H $\beta$	1329 <sup>+308</sup> <sub>-429</sub>	3094 <sup>+0488</sup> <sub>-0619</sub>	1490 <sup>+04</sup> <sub>-04</sub>	3250 <sup>+19</sup> <sub>-18</sub>
	H $\gamma$	1163 <sup>+215</sup> <sub>-364</sub>	2311 <sup>+0814</sup> <sub>-0294</sub>	1148 <sup>+05</sup> <sub>-05</sub>	3648 <sup>+18</sup> <sub>-18</sub>
	He II $\lambda$ 4686	1997 <sup>+158</sup> <sub>-223</sub>	4172 <sup>+1106</sup> <sub>-1083</sub>	1554 <sup>+58</sup> <sub>-70</sub>	4140 <sup>+64</sup> <sub>-64</sub>

**Note.** Columns 3 and 4 give the rms linewidth and FWHM in the rms spectrum. Columns 5 and 6 give the same but in the mean spectrum. All values are corrected for instrumental broadening and the smoothing in Section 2.2.2 (see Section 4) and are reported in the rest frame. The uncertainties only represent the statistical errors—blending, continuum interpolation, and the choice of wavelength windows introduce additional systematic uncertainties (especially for He II  $\lambda$ 4686).

Grier et al. 2013a). We use the most recent calibration by Woo et al. (2015) of  $\langle f \rangle = 4.47 \pm 1.25$  with a scatter of  $0.43 \pm 0.03$  dex (a factor of 2.7). Finally, it is convenient to define the virial product,  $\sigma_L^2 c \tau / G$ , which is an observed quantity that is independent of the mass calibration.

We calculate the statistical uncertainties on the virial products through standard error propagation. As discussed above, there are significant systematic uncertainties on both the linewidths and the lags, which probably dominate the final error budget (see also Section 2.4). We estimate the systematic uncertainty using repeat RM measurements gathered from the literature. There are 17 H $\beta$ -based measurements of the virial product in NGC 5548 over the last 30 years (see Bentz & Katz 2015). The (log) standard deviation of these measurements is 0.16 dex, while the mean statistical uncertainty is 0.10 dex. Taking  $\sigma_{\text{sys}}^2 = \sigma_{\text{rms}}^2 - \sigma_{\text{stat}}^2$ , we estimate a systematic uncertainty floor of 0.13 dex. Experimentation with alternative line windows, continuum interpolations, and detrending procedures suggests that this value (a factor of about 1.3) captures most of the variation in the virial products of our sample. We therefore adopt 0.13 dex as our estimate of the systematic uncertainty on each virial product and add this value in quadrature to the statistical uncertainties for the virial products. For our final mass estimates, we also add in quadrature the the uncertainty in

the mean value of  $\langle f \rangle$  ( $\sim 0.12$  dex) and its intrinsic scatter (0.43 dex). The virial products, final masses, and total uncertainties are given in Table 17.

We discuss the consistency of virial products for the same object derived from different emission lines in Section 5.2, and we comment on the H $\beta$ -derived masses of individual objects below.

- i. MCG+08-11-011 is our most variable object. The black hole mass estimate is  $\sim 2.8 \times 10^7 M_\odot$ , and the uncertainty is dominated by uncertainty in the virial factor  $f$ . Bianchi et al. (2010) found evidence for a relativistically broadened Fe K $\alpha$  line in the X-ray spectrum of this object, but the available mass estimates at that time were uncertain by an order of magnitude ( $10^7$ – $10^8 M_\odot$ ). The results presented here may help measure the spin of the black hole in future studies.
- ii. The mass reported here for NGC 2617 of  $\sim 3.2 \times 10^7 M_\odot$  is in good agreement with the single-epoch mass estimated by Shappee et al. (2014) of  $(4 \pm 1) \times 10^7 M_\odot$ , also using the H $\beta$  emission line. NGC 2617 is the second “changing look” AGN with a direct RM mass measurement. The other object is Mrk 590, which was observed to change from a Seyfert 1.5 to 1.0 to 1.9 over several decades (Denney et al. 2014) and has an RM mass of  $\sim 5 \times 10^7 M_\odot$  (Peterson et al. 2004). In terms of their black hole masses, there is nothing extraordinary about either NGC 2617 or Mrk 590. Our luminosity-independent RM mass also allows us to estimate a more robust Eddington ratio ( $\dot{m}_{\text{Edd}} = L_{\text{Bol}}/L_{\text{Edd}}$ ) than from the single-epoch mass. Assuming a bolometric correction of 10 for the 5100 Å continuum luminosity, we find that  $\dot{m}_{\text{Edd}} = 0.01$ , after correcting for host-galaxy starlight (see Section 5.1). This value is somewhat low, though not atypical, for Seyfert 1 galaxies.
- iii. For NGC 4051, our measurement of the H $\beta$  lag ( $2.24 \pm 0.33$  days) is in good agreement with the estimate of  $1.87 \pm 0.52$  days by Denney et al. (2009b). The measurement is challenging because of the low-amplitude continuum variations, variable host-galaxy contamination from aperture effects (Peterson et al. 1995), and a secular trend in the line light curve.

Our estimate of the virial product  $\sim 1.1 \times 10^5 M_\odot$  is also consistent at the  $2\sigma$  level with the estimate of  $(3.0 \pm 1.0) \times 10^5 M_\odot$  from Denney et al. (2010). The difference is primarily due to a decrease in the linewidth by about  $400 \text{ km s}^{-1}$  compared to the 2007 campaign. The line and continuum wavelength window definitions are somewhat different between the 2014 and 2007 campaigns, and we found that using the wavelength windows from Tables 2 and 3 for the rms spectrum from 2007 reduces the difference to only  $\sim 100 \text{ km s}^{-1}$  (i.e.,  $\sigma_L$  was about 20% larger in 2007 than in 2014). If we use the wavelength regions from Denney et al. (2010), the measurement from 2014 increases by  $\sim 120 \text{ km s}^{-1}$ . This suggests that the virial product is somewhat smaller than that reported by Denney et al. (2010), but the mild  $2\sigma$  discrepancy indicates that the systematic uncertainties are comparable to the formal uncertainties. The remaining  $100$ – $300 \text{ km s}^{-1}$  difference is physical—comparing the rms line profiles between the two campaigns, we found that the core of the H $\beta$  line is much more variable in 2014 than it was in 2007, weighting  $\sigma_L$  to smaller values. The



**Table 17**  
Black Hole Masses

Object	Line	$\tau$ (days)	$\sigma_L$ (km s <sup>-1</sup> )	$\log(c\tau\sigma_V^2/G)^a$ ( $M_\odot$ )	$\log M_{\text{BH}}^b$ ( $M_\odot$ )
(1)	(2)	(3)	(4)	(5)	(6)
MCG+08-11-011	H $\beta$	$14.98^{+0.34}_{-0.28}$	$1466^{+102}_{-174}$	$6.80 \pm 0.15$	$7.45 \pm 0.47$
	H $\gamma$	$12.38^{+0.46}_{-0.49}$	$1604^{+83}_{-82}$	$6.79 \pm 0.14$	$7.44 \pm 0.47$
	He II $\lambda$ 4686	$1.21^{+0.29}_{-0.33}$	$2453^{+125}_{-130}$	$6.20 \pm 0.18$	$6.80 \pm 0.48$
NGC 2617	H $\beta$	$6.38^{+0.44}_{-0.50}$	$2424^{+91}_{-86}$	$6.86 \pm 0.14$	$7.51 \pm 0.47$
	H $\gamma$	$0.81^{+0.59}_{-0.61}$	$3084^{+86}_{-90}$	$6.17 \pm 0.44$	$6.82 \pm 0.63$
	He II $\lambda$ 4686	$1.75^{+0.34}_{-0.38}$	$2020^{+329}_{-572}$	$6.14 \pm 0.26$	$6.79 \pm 0.52$
NGC 4051	H $\beta$	$2.24^{+0.39}_{-0.28}$	$493^{+34}_{-36}$	$5.02 \pm 0.16$	$5.67 \pm 0.47$
	H $\gamma$	$2.40^{+0.86}_{-0.73}$	$641^{+55}_{-59}$	$5.28 \pm 0.21$	$5.93 \pm 0.50$
3C 382	H $\beta$	$52.07^{+3.18}_{-9.46}$	$4552^{+214}_{-163}$	$8.33 \pm 0.14$	$8.98 \pm 0.47$
Mrk 374	H $\beta$	$13.73^{+1.06}_{-1.02}$	$1329^{+308}_{-429}$	$6.67 \pm 0.31$	$7.32 \pm 0.54$
	H $\gamma$	$13.37^{+2.11}_{-2.08}$	$1163^{+215}_{-364}$	$6.55 \pm 0.28$	$7.20 \pm 0.53$

**Notes.** Column 3 gives the adopted lag and its statistical uncertainty,  $\tau_{\text{multi}}$ , from Table 15. Column 4 gives the rms linewidth  $\sigma_L$  from Table 16 of the line profile in the rms residual spectrum and its statistical uncertainty (see Sections 2.4 and 4), corrected to the rest frame. Column 5 gives the virial product, which is independent of any calibration to the  $M_{\text{BH}}-\sigma_*$  relation. Column 6 gives the SMBH mass using the  $M_{\text{BH}}-\sigma_*$  calibration from Woo et al. (2015) with  $f = 4.47 \pm 1.25$ .

<sup>a</sup> Includes a 0.13 dex systematic uncertainty, added in quadrature to the statistical uncertainties propagated from Columns 3 and 4.

<sup>b</sup> Include uncertainty in the mean value of  $f$  (0.12 dex) and its intrinsic scatter (0.43 dex) added in quadrature to the uncertainties from Column 5.

lag has only increased by 0.26 days (19%), so the virial product shows a net decrease. This might indicate a change in the geometry and/or dynamics of the BLR. The dynamical time is of order only two or three years at two light days from a  $10^6 M_\odot$  black hole, so such a change cannot be ruled out a priori. A comparison of the velocity-resolved reverberation signals between 2007 and 2014 is therefore especially interesting.

Our SMBH mass estimate of  $\sim 4.7 \times 10^5 M_\odot$  for NGC 4051 is at the very low end of the SMBH scale, and there are only two other RM masses below  $10^6 M_\odot$ : NGC 4395 (Peterson et al. 2005; Edri et al. 2012) and UGC 06728 (Bentz et al. 2016).

- iv. In 3C 382, the black hole mass is about  $9.6 \times 10^8 M_\odot$ , and a large source of uncertainty is the H $\beta$  lag. The  $\sim 52$  day lag is driven by the gentle inflection in the line light curve observed near the middle of the spectroscopic campaign, which was also observed in the imaging data about one month before the MDM observations began. The uncertainties on the H $\beta$  line lag are therefore quite large. By RM standards, 3C 382 is also at a moderate redshift ( $z \sim 0.06$ ) and faint ( $V \sim 15.4$ ), putting it near the limit of feasibility for monitoring campaigns with a 1m-class telescope.

Several estimates of the BLR orientation exist for this object. Emission from the radio lobes in 3C 382 dominates over that of the core, indicating that the system is viewed more edge on (Wills & Browne 1986 give the core-to-lobe ratio as  $\sim 0.1$ ). However, Eracleous et al. (1995) found an inclination of  $45^\circ$  from dynamical modeling of the double-peaked broad H $\alpha$  line and show that this estimate is consistent with the radio properties. Velocity-delay maps and dynamical modeling of this object would be an interesting test of this inclination measurement. Unfortunately, the width of the continuum

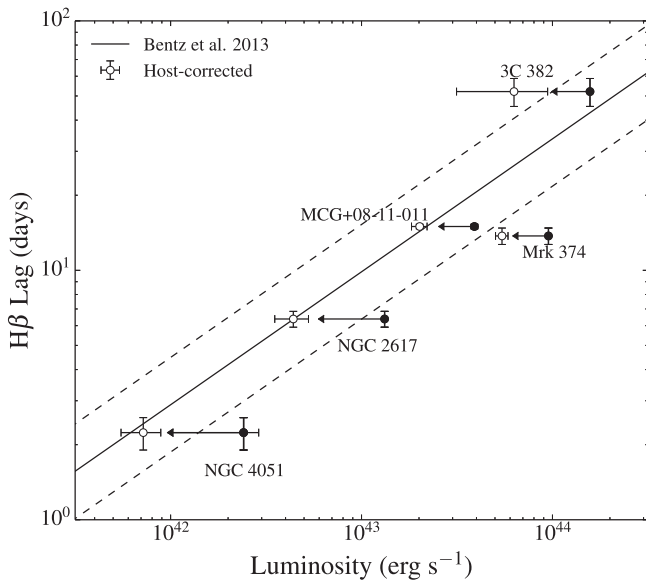
autocorrelation function and the low S/N of the line light curves are poorly suited for these experiments. On the other hand, a moderately inclined disk is broadly consistent with the double-peaked rms H $\beta$  and H $\gamma$  line profiles, and velocity-binned mean time delays may still provide interesting constraints on the BLR structure.

- v. Mrk 374 is our least variable source. Although the H $\beta$  lag is detected at a statistically significant level, the uncertainty on the ICCF centroid is somewhat larger than for the other objects ( $\sim 33\%$ ). The mass is  $\sim 2.09 \times 10^7 M_\odot$  and the dominant uncertainty is from the linewidth measurement—it is clear from Figure 6 that the variability of the lines is very small and that there is some ambiguity in where the line profile begins and ends. At a redshift of  $\sim 0.04$ , Mrk 374 is one of our fainter sources ( $V = 15.0$  mag), and, similar to 3C 382, it is near the practical limits of a monitoring campaign lead by a 1m-class telescope.

## 5. Discussion

### 5.1. Radius–Luminosity Relation

Figure 10 shows the H $\beta$  lags of our five objects as a function of luminosity, the so-called radius–luminosity ( $R-L$ ) relation (Kaspi et al. 2000, 2005; Bentz et al. 2009, 2013). To estimate the luminosities, we first take the mean of the 5100 Å light curve and correct for Galactic extinction using the extinction map of Schlafly & Finkbeiner (2011) and a Cardelli et al. (1989) extinction law with  $R_V = 3.1$ . We then convert the flux to luminosity using the luminosity distances in Table 1. In the case of NGC 4051, which has a large peculiar velocity relative to the Hubble flow ( $z \sim 0.002$ ), we use a Tully–Fischer distance of 17.1 Mpc (Tully et al. 2008). This distance is uncertain by about 20%, and improving this measurement is an important step to investigate any discrepancies of this object



**Figure 10.** Radius–luminosity relation for the targets of this study, compared to the relation from Bentz et al. (2013). Luminosities are estimated from the mean of the continuum light curves corrected for Galactic extinction. The solid black line shows the best-fit relation measured by Bentz et al. (2013) and the dashed black lines show the dispersion around the best fit. Open circles show the luminosities corrected for host-galaxy starlight, which results in excellent agreement with the relation from Bentz et al. (2013).

from the  $R$ – $L$  relation and to estimate its true Eddington ratio. For these purposes, an *HST* program has recently been approved to obtain a Cepheid distance to NGC 4051 (*HST* GO-14697; PI Peterson).

The final values of  $\lambda L_{5100\text{Å}}$  are reported in Table 1, along with the adopted Galactic values of  $E(B - V)$ . We find that our objects all lie close to, but slightly below (except for 3C 382), the  $R$ – $L$  relation. The major systematic uncertainties are internal extinction in the AGN and host-galaxy contamination. Internal extinction may move the points farther from the  $R$ – $L$  relation, but this effect is expected to be small. On the other hand, host-galaxy contamination can be very significant, especially for low-luminosity objects.

In order to correct for host contamination, we model high-resolution images of the targets and isolate the host-galaxy flux. This has previously been done for NGC 4051 (Bentz et al. 2006, 2013), and MCG+08-11-011, NGC 2617, and Mrk 374 were recently observed with *HST* for this purpose (*HST* GO-13816; PI Bentz). We also retrieved archival optical WFPC2 imaging of 3C 382 (*HST* GO-6967, PI Sparks), but the data are not ideal for image decompositions and we discuss the host-galaxy flux estimate for this object separately. A more detailed analysis of the *HST* GO-13816 data and image decompositions will be presented in future work (M. C. Bentz et al, in preparation). However, following the procedures described by Bentz et al. (2013), we made preliminary estimates of the host-galaxy contributions in the MDM aperture ( $15''0 \times 5''0$  aligned at position angle  $0^\circ$ ). The results are given in Table 1 (uncertainties on these values are estimated at 10% and included in Figure 10). Applying this correction shows that host contamination accounts for the entire discrepancy between the observed luminosities and the  $R$ – $L$  relation. The largest deviation from the  $R$ – $L$  relation is Mrk 374, but the offset is only slightly greater than the  $1\sigma$  scatter of the relation.

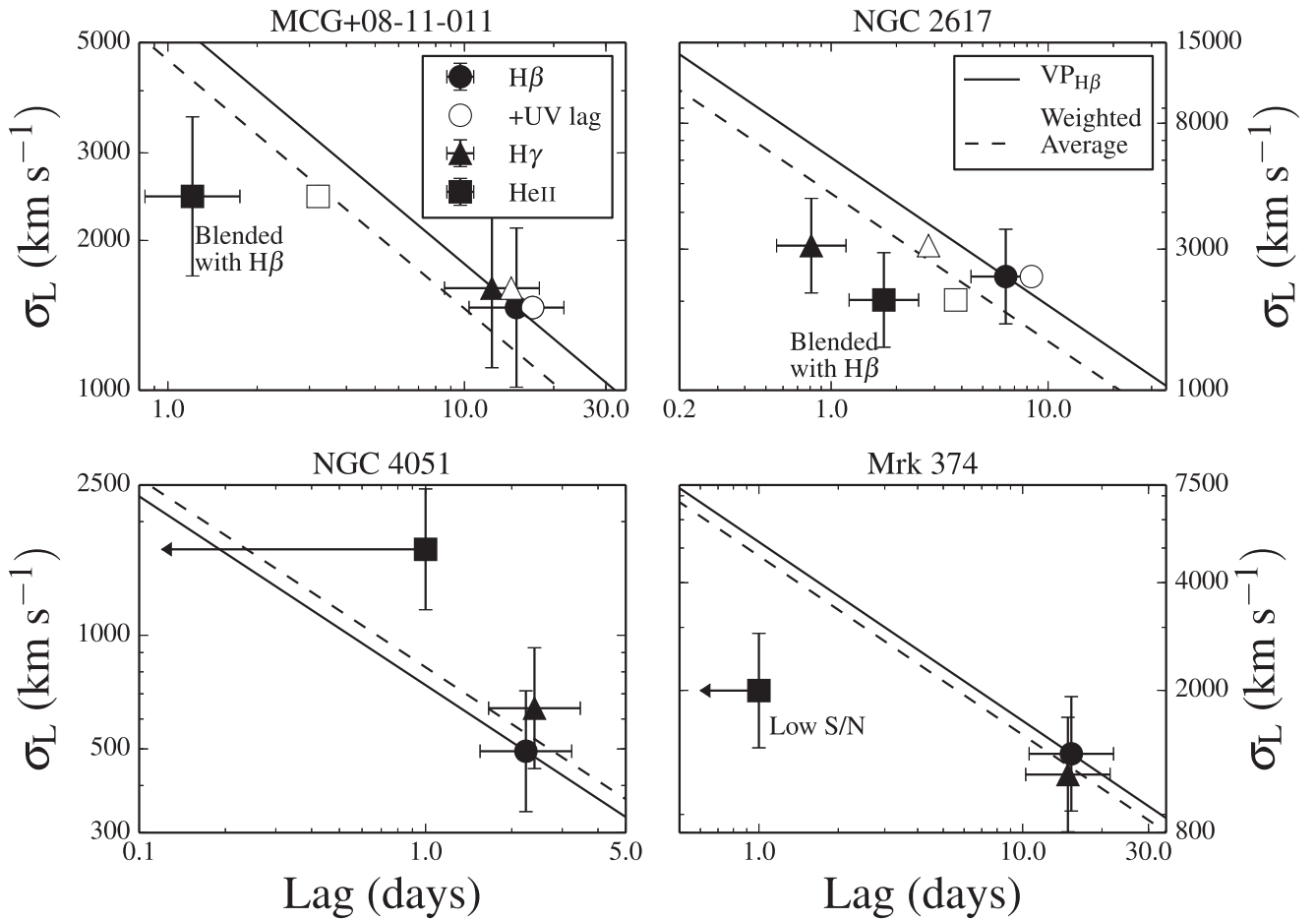
3C 382 resides in a giant elliptical galaxy and there may be a significant contribution from the host’s starlight—several stellar absorption features are visible in the mean spectrum in Figure 5. In the archival *HST* images, the galaxy nucleus is saturated, hindering our ability to robustly remove the AGN flux and isolate the host’s starlight. The main problem is that the Sérsic index of the host galaxy is degenerate with the saturated core and tends to drift toward unreasonably high values ( $n \approx 7.6$ ) when fitting the image in the same way as Bentz et al. (2013). Fixing the Sérsic index to more typical values (between 2 and 4) leads to host fluxes in the MDM aperture between  $2.2$  and  $2.7 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Å}^{-1}$ , about 77% of the observed luminosity ( $\log \lambda L_{\text{host}} = 44.04$  to  $44.12$  [ $\text{erg s}^{-1}$ ], after correcting for Galactic extinction). This estimate can be checked using the equivalent width (EW) of the prominent Mg absorption feature at  $5200 \text{ Å}$  rest frame ( $5460 \text{ Å}$  observed frame). In our mean spectrum, we find an EW of  $2.8 \text{ Å}$ . In typical elliptical galaxy spectra, we find the EW is about  $6.7$ – $7.3 \text{ Å}$ , depending on the continuum estimation and assumptions about the host-galaxy properties.<sup>50</sup> This implies that the featureless AGN continuum dilutes the absorption feature by a factor of  $2.4$ – $2.6$ , so that the host galaxy contributes approximately 40% of the observed luminosity. This rough estimate is a factor of two lower than the result from image decomposition, but the two values span the range of host-contributions from the other objects in our sample (42%–71% of the observed luminosity, see Table 1). We therefore adopt a host correction of  $(60 \pm 20)\%$  of the observed luminosity ( $\log \lambda L_{5100\text{Å}} = 43.98 \pm 0.15$  [ $\text{erg s}^{-1}$ ]), and we note that this estimate can easily be improved by obtaining unsaturated high-resolution images. The host correction moves 3C 382 away from the  $R$ – $L$  relation, just beyond the  $1\sigma$  dispersion. However, considering the large uncertainties, there does not appear to be any evidence that 3C 382 has an anomalous  $H\beta$  lag for its luminosity.

## 5.2. Virialization of the BLR

With the measurement of BLR velocity dispersions at a range of radii, it is possible to test if the BLR is virialized. Virialized dynamics predict  $V(r) \propto r^{-1/2}$ , where the constant of proportionality depends on the SMBH mass and BLR inclination/kinematics. If the BLR is virialized, the virial products  $\sigma_L^2 c \tau / G$  derived from different line species should be consistent with each other, assuming similar geometries and dynamics for the line-emitting gas.

In Table 17, the maximum differences between  $\log \sigma_L^2 c \tau / G$  for each object are  $3.3\sigma$  in MCG+08-11-011,  $2.8\sigma$  in NGC 2617,  $1.2\sigma$  in NGC 4051, and  $0.4\sigma$  in Mrk 374. For NGC 4051 and Mrk 374, these differences are not significant. For MCG+08-11-011 and NGC 2617, the  $H\beta$  and He II  $\lambda 4686$  virial products are marginally discrepant at about  $2.5$ – $3.3\sigma$ . We show these results Figure 11, which displays the linewidths  $\sigma_L$  as a function of lag  $\tau_{\text{multi}}$ , and the relation  $\sigma_L \propto \tau_{\text{multi}}^{-1/2}$  normalized by the value for  $H\beta$ . In this figure, we have applied a  $0.13$  dex uncertainty to both the lag  $\tau$  and linewidth  $\sigma_L$ , representative of the characteristic systematic uncertainties.

<sup>50</sup> We used two different templates for the “standard” giant elliptical spectrum: observations of the E0 galaxy NGC 1407 used to construct empirical templates (Kinney et al. 1996; Denney et al. 2009a), and a synthetic stellar population model from Bruzual & Charlot (2003) consisting of a single 11 Gyr population at solar metallicity.



**Figure 11.** Linewidth  $\sigma_L$  as a function of lag  $\tau_{\text{multi}}$  from Columns 3 and 4 of Table 17. The solid lines show the virial relation  $\sigma_L \propto \tau_{\text{multi}}^{-1/2}$  normalized by the  $\text{H}\beta$  virial product ( $\text{VP}_{\text{H}\beta}$ ). The dashed lines show the virial relation using a statistical uncertainty-weighted average of the different emission-line virial products. The open points for MCG+08-11-011 and NGC 2617 show the effect of adding a hypothetical 2.0 day UV-optical lag, similar to that found in NGC 5548 (see Section 5.2). For NGC 4051 and Mrk 374, the  $\text{He II } \lambda 4686$  lags are consistent with zero, and we show these as upper limits at 1 day. Uncertainties of 0.13 dex are assigned to both  $\tau_{\text{multi}}$  and  $\sigma_L$ , to approximately represent the level of systematic uncertainty associated with the virial products.

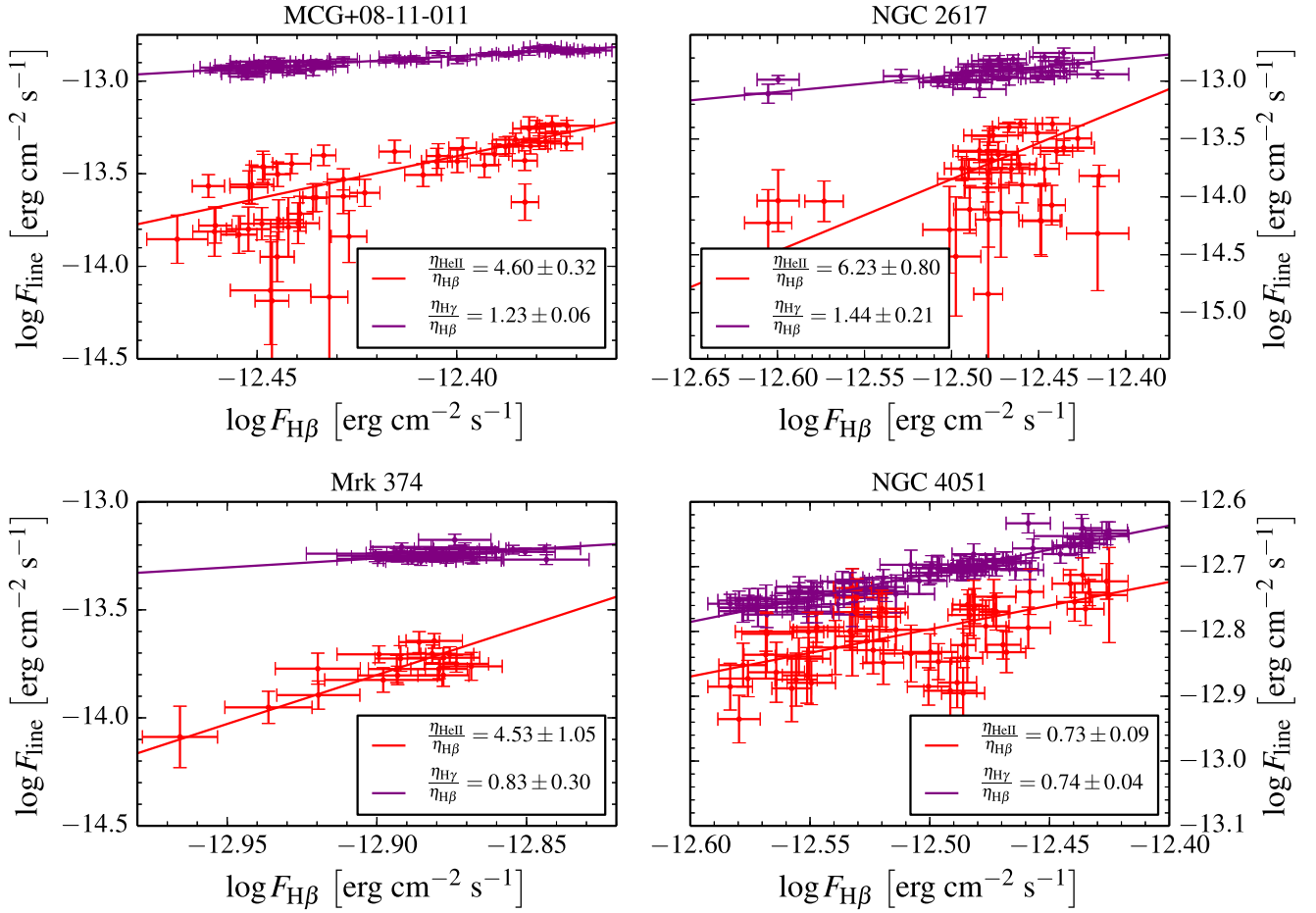
While the  $\text{H}\gamma$  points generally agree with the  $\text{H}\beta$  relation, the  $\text{He II } \lambda 4686$  points have very large offsets.

There are many systematic issues that could account for these differences. As discussed in Section 4, the red wing of  $\text{He II } \lambda 4686$  is blended with the blue wing of  $\text{H}\beta$  in both MCG+08-11-011 and NGC 2617. The  $\text{He II } \lambda 4686$  velocity is therefore likely underestimated because we cannot follow its red wing underneath  $\text{H}\beta$ . The  $\text{He II } \lambda 4686$  lags are also small compared to the monitoring cadence, and the lag is only marginally detected at  $2.3\sigma$  in NGC 2617. Furthermore, the choice of line window and continuum interpolation can have a significant effect on the lag and linewidths. Finally, we must assume that the 5100 Å continuum light curve is a suitable proxy for the ionizing flux variations at extreme UV wavelengths. In NGC 5548, we found a  $\sim 2$ -day lag between the far-UV and optical emission (Edelson et al. 2015; Fausnaugh et al. 2016). If a similar lag exists in these objects, it would change the  $\text{He II } \lambda 4686$  virial products by a significant amount (0.3–0.4 dex), while the change in the  $\text{H}\beta$  virial products would be much smaller (0.05–0.11 dex). The effect of adding a 2-day UV-optical lag to the optical-line lags is shown in Figure 11, and the additional lag would reduce the discrepancies in the virial products to  $1.3\sigma$  for MCG+08-11-011 and  $2.0\sigma$  for NGC 2617. These AGNs have masses and luminosities similar to NGC 5548, so the existence of a

UV-optical lag of this magnitude is very likely. Although a UV-optical lag affects the virial product and the characteristic size of the BLR, it does not affect the final mass estimate because the virial factor  $\langle f \rangle$  is calibrated using the  $M_{\text{BH}}-\sigma_*$  relation (see Fausnaugh et al. 2016; Pei et al. 2017).

If the remaining discrepancies are real, they indicate different dynamics and geometries for the  $\text{He II } \lambda 4686$  line-emitting gas compared to that of  $\text{H}\beta$ . This might be plausible, since  $\text{He II } \lambda 4686$  is a high-ionization state line and may originate in very different physical conditions than the Balmer lines (for example, a disk wind). If  $\text{He II } \lambda 4686$  has different dynamics than  $\text{H}\beta$ , it would be necessary to calibrate a different virial factor  $\langle f \rangle$  for the  $\text{He II } \lambda 4686$  line when calculating the SMBH masses. However, we cannot rule out systematic effects and it is unclear if the  $\text{He II } \lambda 4686$  discrepancies are physical. If systematic issues do account for the discrepancies, then the dynamics of the BLRs in these AGNs would be consistent with virialized motion, as has been found for other AGNs (Peterson et al. 2004).

The  $\text{H}\beta$  light curves and line profiles have much higher S/N and very clear lags compared to both  $\text{He II } \lambda 4686$  and  $\text{H}\gamma$ , resulting in more reliable black hole masses. If we combine the virial products in Table 17 using statistical uncertainty-weighted average, the virial relation changes little, as shown in



**Figure 12.** Responsivities of optical recombination lines relative to  $H\beta$ . The results for  $H\gamma$  are shown in purple and those for  $He\ II\ \lambda 4686$  are shown in red. Solid lines show weighted linear least-squares fits (accounting for uncertainties in both coordinates)—the slope of the fit gives the relative responsivity, which is listed in each panel. See Section 5.3 for more details.

Figure 11 with the dashed lines. We therefore take the  $H\beta$  masses for our standard SMBH mass estimates.

### 5.3. Photoionization Physics

Photoionization models make predictions about the structure of the BLR that can be tested with RM of multiple recombination lines. The locally optimally emitting cloud model (Baldwin et al. 1995) provides a natural explanation for the general similarity of AGN spectra and predicts radial stratification of the BLR—high-ionization state lines, such as  $He\ II\ \lambda 1640/4686$  and  $C\ IV\ \lambda 1549$ , should be primarily emitted at smaller radii than low-ionization state lines such as  $H\beta$  and  $Mg\ II\ \lambda 2798$ . Korista & Goad (2004, hereinafter KG04) used this model to predict that the responsivity of high-order Balmer lines should be greater than that of low-order lines (in the sense that  $H\gamma > H\beta > H\alpha$ ). KG04 also predicted that high-ionization state lines such as  $He\ II\ \lambda 4686$  should have greater responsivity than all of the Balmer lines. Radial stratification of the BLR in NGC 5548 was first observed by Clavel et al. (1991), and has since been observed in several other objects (Peterson et al. 2004; Grier et al. 2013b). In addition, the expected trends of responsivity with ionization state/species have been confirmed in 16 AGNs by LAMP (Bentz et al. 2010; Barth et al. 2015).

We confirm these results for the four objects with multiple line light curves presented here. The  $He\ II\ \lambda 4686$  lags in MCG

+08-11-011 and NGC 2617 are less than 2 days, while the  $H\beta$  lags are 14.82 and 6.38 days, respectively, clearly indicating radial stratification. Furthermore, the fractional variability of the light curves, as measured by  $F_{var}$  (Table 4), is generally larger for  $H\gamma$  than  $H\beta$  (or comparable for NGC 4051 and Mrk 374), while  $F_{var}$  for  $He\ II\ \lambda 4686$  is always much greater than for the Balmer lines (although it is only slightly higher in NGC 4051). This implies that the relative line responsivities are  $He\ II\ \lambda 4686 \gg H\gamma > H\beta$ , in agreement with the photoionization models. We also find that the  $H\gamma$  lags are slightly shorter than the  $H\beta$  lags within the same object (except for NGC 4051). KG04 showed that shorter lags are a natural consequence of the higher responsivity of  $H\gamma$  compared to  $H\beta$ .

The formal definition of the responsivity of an emission line is

$$\eta_{line} = \frac{\Delta \log F_{line}}{\Delta \log \Phi}, \quad (10)$$

where  $F_{line}$  is the line flux and  $\Phi$  is the photoionizing flux (KG04). The parameter  $\eta_{line}$  is therefore a measure of how efficiently the BLR converts a change in the photoionizing flux into a change in line emission. The ionizing flux  $\Phi$  cannot be observed directly because these photons are at far-UV wavelengths ( $< 912\ \text{\AA}$ ). Therefore, we cannot measure  $\eta_{line}$  directly, but we can measure the relative responsivity  $\eta_{line1}/\eta_{line2} = \Delta \log F_{line1}/\Delta \log F_{line2}$ .



We present rough measurements of the relative responsivity of  $H\beta$ ,  $H\gamma$ , and  $\text{He II } \lambda 4686$  in Figure 12. For each object, we first removed the lags of each line from the corresponding light curve. We then matched observed points to the nearest day between the  $H\beta$  light curves and  $H\gamma$  or  $\text{He II } \lambda 4686$  light curves. The ratio  $\eta_{\text{line}}/\eta_{H\beta}$  then corresponds to the slope of a linear least-squares fit to the data in the  $\log F_{H\beta}$ – $\log F_{\text{line}}$  plane.

We find that  $\eta_{H\gamma}/\eta_{H\beta}$  ranges from 0.74 to 1.44 and that  $\eta_{\text{He II}}/\eta_{H\beta}$  ranges between 0.73 and 6.23. NGC 4051, with  $\eta_{\text{He II}}/\eta_{H\beta} \sim 0.73$ , is an outlier, probably caused by over-subtracting the continuum before integrating the line flux. For comparison, KG04 calculated  $\eta_{\text{line}}$  for a fiducial model of the BLR in NGC 5548, which includes an empirically motivated but ad hoc parameterization of the ionizing flux. From their Table 1,  $\eta_{H\gamma}/\eta_{H\beta}$  ranges between 1.03 and 1.07, depending on the flux state of the AGN, while  $\eta_{\text{He II}}/\eta_{H\beta}$  ranges from 1.26 to 1.61. Thus, while our fits for  $\eta_{H\gamma}/\eta_{H\beta}$  are in reasonable agreement with this fiducial model, the values of  $\eta_{\text{He II}}/\eta_{H\beta}$  are much larger than the model’s prediction. The spread of  $\eta_{\text{line}}/\eta_{H\beta}$  in our fits is also fairly large, which may indicate a diversity of photoionization conditions in the BLRs of different objects (perhaps due to harder or softer ionizing fluxes than assumed for NGC 5548).

Our estimates of the relative responsivities are sensitive to the total flux of the line light curves. For example, the sublinear slopes for  $\eta_{H\gamma}/\eta_{H\beta}$  in NGC 4051 and Mrk 374 could be explained by missing variable line flux, perhaps in the wings of the line during low-flux states, or excess constant flux from the narrow emission lines or host-galaxy starlight. On the other hand, large values of  $\eta_{\text{He II}}/\eta_{H\beta}$  might be explained by contamination by Fe II lines or misestimation of the continuum.

## 6. Summary and Future Prospects

We have presented the initial analysis of data from an intensive RM monitoring campaign carried out in the first half of 2014. We succeeded in measuring continuum-line lags for six targets, five of which are presented here. (For NGC 5548, see Pei et al. 2017.) Our main results are as follows.

- i. Four new SMBH masses, as well as a refined measurement for NGC 4051.
- ii. In addition to measuring  $H\beta$  lags for all five targets, we measure  $H\gamma$  lags in four objects and  $\text{He II } \lambda 4686$  lags in two objects.
- iii. Using the  $\text{He II } \lambda 4686$  lags (or their upper limits), we show that the BLR is radially stratified. Although the  $\text{He II } \lambda 4686$  virial products are somewhat smaller than those derived from  $H\beta$ , systematic effects such as blending in the line wings and the choice of continuum interpolation may account for these discrepancies. The BLRs are otherwise consistent with virialized dynamics with  $V(r) \propto r^{-1/2}$ .
- iv. We find that  $\text{He II } \lambda 4686$  is more responsive than the Balmer lines, and that  $H\gamma$  is more responsive than  $H\beta$ , in agreement with predictions from photoionization modeling.

Many modern RM experiments are focused on measuring velocity-resolved reverberation signatures, in order to investigate the geometry and dynamics of the BLR. There are only six AGNs with published velocity-delay maps (Ulrich & Horne 1996; Bentz et al. 2010; Grier et al. 2013b) and

five AGNs with direct BLR dynamical models (Pancoast et al. 2014b; one AGN, Arp 151, has both). The data presented here are of exceptional quality and very well-calibrated—based on the cadence and S/N of these observations, we have an excellent prospect of recovering velocity-delay maps and dynamical models in three objects (MCG+08-11-011, NGC 2617, and NGC 4051). This will expand the sample of AGNs with detailed BLR information by  $\sim 30\%$ , demonstrating the continuing importance of targeted and intensive monitoring campaigns.

M.M.F. acknowledges financial support from a Presidential Fellowship awarded by The Ohio State University Graduate School. NSF grant AST-1008882 supported M.M.F., G.D.R., B.M.P., and R.W.P.; M.C.B. gratefully acknowledges support through NSF CAREER grant AST-1253702 to Georgia State University. K.D.D. is supported by an NSF AAPF fellowship awarded under NSF grant AST-1302093. C.S.K. is supported by NSF grant AST-1515876. K.H. acknowledges support from STFC grant ST/M001296/1. This material is based in part upon work supported by the National Science Foundation (NSF) Graduate Research Fellowship Program under grant No. DGE-0822215, awarded to C.B.H.; A.M.M. acknowledges the support of NSF grant AST-1211146. M.E. thanks the members of the Center for Relativistic Astrophysics at Georgia Tech, where he was based during the observing campaign, for their warm hospitality. J.S.S. acknowledges CNPq, National Council for Scientific and Technological Development, Brazil. J.T. acknowledges support from NSF grant AST-1411685. Work by S.V. Jr is supported by the National Science Foundation Graduate Research Fellowship under grant No. DGE-1343012. Work by W.Z. was supported by NSF grant AST-1516842. T.W.-S.H. is supported by the DOE Computational Science Graduate Fellowship, grant No. DE-FG02-97ER25308. E.R.C. and S.M.C. gratefully acknowledge the receipt of research grants from the National Research Foundation (NRF) of South Africa. T.T. acknowledges support by the National Science Foundation through grant AST-1412315 “Collaborative Research: New Frontiers in Reverberation Mapping,” and by the Packard Foundation through a Packard Research Fellowship. D.J.S. acknowledges support from NSF grants AST-1412504 and AST-1517649. A.J.B. and L.P. have been supported by NSF grant AST-1412693. B.J.S. is supported by NASA through Hubble Fellowship grant HF-51348.001 awarded by the Space Telescope Science Institute that is operated by the Association of Universities for Research in Astronomy, Inc., for NASA, under contract NAS 5-26555.

This work makes use of observations with the NASA/ESA *Hubble Space Telescope*. M.C.B. acknowledges support through grant *HST* GO-13816 from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555. This work is based on observations obtained at the MDM Observatory, operated by Dartmouth College, Columbia University, Ohio State University, Ohio University, and the University of Michigan. This paper is partly based on observations collected at the Wise Observatory with the C18 telescope. The C18 telescope and most of its equipment were acquired with a grant from the Israel Space Agency (ISA) to operate a Near-Earth Asteroid Knowledge Center at Tel Aviv University. The Fountainwood Observatory would like to thank the HHMI for its support of science

research for undergraduate students at Southwestern University. This research has made use of NASA's Astrophysics Data System, as well as the NASA/IPAC Extragalactic Database (NED), which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

**Facilities:** McGraw-Hill, *HST*, Wise Observatory, Fountainwood Observatory BYU:0.9m, CrAO:0.7m, WIRO, LCOGT, SSO:1m.

**Software:** Astropy (Astropy Collaboration et al. 2013), IRAF (Tody 1986), ISIS (Alard & Lupton 1998), JAVELIN (Zu et al. 2011), mapspec (Fausnaugh 2017), Matplotlib (Hunter 2007), Numpy (van der Walt et al. 2011), Scipy (Oliphant 2007).

## References

- Alard, C., & Lupton, R. H. 1998, *ApJ*, **503**, 325
- Antonucci, R. 1993, *ARA&A*, **31**, 473
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33
- Balbus, S. A., & Hawley, J. F. 1998, *RvMP*, **70**, 1
- Baldwin, J., Ferland, G., Korista, K., & Verner, D. 1995, *ApJL*, **455**, L119
- Barth, A. J., Bennert, V. N., Canalizo, G., et al. 2015, *ApJS*, **217**, 26
- Bentz, M. C., Batista, M., Seals, J., et al. 2016, *ApJ*, **831**, 2
- Bentz, M. C., Denney, K. D., Grier, C. J., et al. 2013, *ApJ*, **767**, 149
- Bentz, M. C., Horne, K., Barth, A. J., et al. 2010, *ApJL*, **720**, L46
- Bentz, M. C., & Katz, S. 2015, *PASP*, **127**, 67
- Bentz, M. C., Peterson, B. M., Netzer, H., Pogge, R. W., & Vestergaard, M. 2009, *ApJ*, **697**, 160
- Bentz, M. C., Peterson, B. M., Pogge, R. W., Vestergaard, M., & Onken, C. A. 2006, *ApJ*, **644**, 133
- Bianchi, S., de Angelis, I., Matt, G., et al. 2010, *A&A*, **522**, A64
- Blackburne, J. A., Pooley, D., Rappaport, S., & Schechter, P. L. 2011, *ApJ*, **729**, 34
- Blandford, R. D., & McKee, C. F. 1982, *ApJ*, **255**, 419
- Brosch, N., Polishook, D., Shporer, A., et al. 2008, *Ap&SS*, **314**, 163
- Brown, T. M., Baliber, N., Bianco, F. B., et al. 2013, *PASP*, **125**, 1031
- Bruzual, G., & Charlot, S. 2003, *MNRAS*, **344**, 1000
- Burbidge, E. M. 1967, *ARA&A*, **5**, 399
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, **345**, 245
- Clavel, J., Reichert, G. A., Alloin, D., et al. 1991, *ApJ*, **366**, 64
- Clavel, J., Wamsteker, W., & Glass, I. S. 1989, *ApJ*, **337**, 236
- Collin-Souffrin, S. 1987, *A&A*, **179**, 60
- Czerny, B., & Hryniewicz, K. 2011, *A&A*, **525**, L8
- Denney, K. D., De Rosa, G., Croxall, K., et al. 2014, *ApJ*, **796**, 134
- Denney, K. D., Horne, K., Shen, Y., et al. 2016, *ApJS*, **224**, 14
- Denney, K. D., Peterson, B. M., Dietrich, M., Vestergaard, M., & Bentz, M. C. 2009a, *ApJ*, **692**, 246
- Denney, K. D., Peterson, B. M., Pogge, R. W., et al. 2010, *ApJ*, **721**, 715
- Denney, K. D., Pogge, R. W., Assef, R. J., et al. 2013, *ApJ*, **775**, 60
- Denney, K. D., Watson, L. C., Peterson, B. M., et al. 2009b, *ApJ*, **702**, 1353
- De Rosa, G., Peterson, B. M., Ely, J., et al. 2015, *ApJ*, **806**, 128
- Dietrich, M., Peterson, B. M., Grier, C. J., et al. 2012, *ApJ*, **757**, 53
- Du, P., Lu, K.-X., Hu, C., et al. 2016, *ApJ*, **820**, 27
- Edelson, R., Gelbord, J. M., Horne, K., et al. 2015, *ApJ*, **806**, 129
- Edelson, R., Vaughan, S., Malkan, M., et al. 2014, *ApJ*, **795**, 2
- Edri, H., Rafter, S. E., Chelouche, D., Kaspi, S., & Behar, E. 2012, *ApJ*, **756**, 73
- Elitzur, M., & Netzer, H. 2016, *MNRAS*, **459**, 585
- Elitzur, M., & Shlosman, I. 2006, *ApJL*, **648**, L101
- Elvis, M. 2000, *ApJ*, **545**, 63
- Eracleous, M., Livio, M., Halpern, J. P., & Storchi-Bergmann, T. 1995, *ApJ*, **438**, 610
- Fausnaugh, M. M. 2017, *PASP*, **129**, 024007
- Fausnaugh, M. M., Denney, K. D., Barth, A. J., et al. 2016, *ApJ*, **821**, 56
- Fausnaugh, M. M., Starkey, D. A., Horne, K., et al. 2017, *ApJ*, submitted
- Gaskell, C. M., Goosmann, R. W., & Klimek, E. S. 2008, *MmSAI*, **79**, 1090
- Gaskell, C. M., & Peterson, B. M. 1987, *ApJS*, **65**, 1
- Goad, M. R., Korista, K. T., De Rosa, G., et al. 2016, *ApJ*, **824**, 11
- Goad, M. R., Korista, K. T., & Ruff, A. J. 2012, *MNRAS*, **426**, 3086
- Grier, C. J., Martini, P., Watson, L. C., et al. 2013a, *ApJ*, **773**, 90
- Grier, C. J., Peterson, B. M., Horne, K., et al. 2013b, *ApJ*, **764**, 47
- Grier, C. J., Peterson, B. M., Pogge, R. W., et al. 2012, *ApJ*, **755**, 60
- Hartman, J. D., Bakos, G., Stanek, K. Z., & Noyes, R. W. 2004, *AJ*, **128**, 1761
- Higginbottom, N., Proga, D., Knigge, C., et al. 2014, *ApJ*, **789**, 19
- Horne, K. 1994, in ASP Conf. Ser. 69, Reverberation Mapping of the Broad-Line Region in Active Galactic Nuclei, ed. P. M. Gondhalekar, K. Horne, & B. M. Peterson (San Francisco, CA: ASP), 23
- Horne, K., Peterson, B. M., Collier, S. J., & Netzer, H. 2004, *PASP*, **116**, 465
- Hunter, J. D. 2007, *Computing in Science & Engineering*, **9**, 90
- Kasliwal, V. P., Vogeley, M. S., & Richards, G. T. 2015, *MNRAS*, **451**, 4328
- Kaspi, S., Maoz, D., Netzer, H., et al. 2005, *ApJ*, **629**, 61
- Kaspi, S., Smith, P. S., Netzer, H., et al. 2000, *ApJ*, **533**, 631
- Kelly, B. C., Bechtold, J., & Siemiginowska, A. 2009, *ApJ*, **698**, 895
- Kinney, A. L., Calzetti, D., Bohlin, R. C., et al. 1996, *ApJ*, **467**, 38
- Kishimoto, M., Hönig, S. F., Beckert, T., & Weigelt, G. 2007, *A&A*, **476**, 713
- Koratkar, A., & Blaes, O. 1999, *PASP*, **111**, 1
- Korista, K. T., & Goad, M. R. 2004, *ApJ*, **606**, 749 (KG04)
- Kozłowski, S. 2016a, *MNRAS*, **459**, 2787
- Kozłowski, S. 2016b, *ApJ*, **826**, 118
- LaMassa, S. M., Cales, S., Moran, E. C., et al. 2015, *ApJ*, **800**, 144
- Lawrence, A. 1991, *MNRAS*, **252**, 586
- MacLeod, C. L., Ivezić, Ž., Kochanek, C. S., et al. 2010, *ApJ*, **721**, 1014
- MacLeod, C. L., Ross, N. P., Lawrence, A., et al. 2016, *MNRAS*, **457**, 389
- McHardy, I. M., Connolly, S. D., Peterson, B. M., et al. 2016, *AN*, **337**, 500
- McLure, R. J., & Dunlop, J. S. 2004, *MNRAS*, **352**, 1390
- Mejía-Restrepo, J. E., Trakhtenbrot, B., Lira, P., Netzer, H., & Capellupo, D. M. 2016, *MNRAS*, **460**, 187
- Miller, L., Turner, T. J., Reeves, J. N., et al. 2010, *MNRAS*, **403**, 196
- Mor, R., & Netzer, H. 2012, *MNRAS*, **420**, 526
- Morgan, C. W., Kochanek, C. S., Morgan, N. D., & Falco, E. E. 2010, *ApJ*, **712**, 1129
- Mosquera, A. M., Kochanek, C. S., Chen, B., et al. 2013, *ApJ*, **769**, 53
- Murray, N., & Chiang, J. 1997, *ApJ*, **474**, 91
- Neukova, M., Sirocky, M. M., Nikutta, R., Ivezić, Ž., & Elitzur, M. 2008, *ApJ*, **685**, 160
- Netzer, H., & Laor, A. 1993, *ApJL*, **404**, L51
- Oke, J. B. 1990, *AJ*, **99**, 1621
- Oknyanskij, V. L., & Horne, K. 2001, in ASP Conf. Ser. 224, Probing the Physics of Active Galactic Nuclei, ed. B. M. Peterson, R. W. Pogge, & R. S. Polidan (San Francisco, CA: ASP), 149
- Oliphant, T. E. 2007, *Computing in Science & Engineering*, **9**, 10
- Onken, C. A., Ferrarese, L., Merritt, D., et al. 2004, *ApJ*, **615**, 645
- Osterbrock, D. E., Koski, A. T., & Phillips, M. M. 1975, *ApJL*, **197**, L41
- Osterbrock, D. E., Koski, A. T., & Phillips, M. M. 1976, *ApJ*, **206**, 898
- Pancoast, A., Brewer, B. J., & Treu, T. 2014a, *MNRAS*, **445**, 3055
- Pancoast, A., Brewer, B. J., Treu, T., et al. 2014b, *MNRAS*, **445**, 3073
- Park, D., Kelly, B. C., Woo, J.-H., & Treu, T. 2012a, *ApJS*, **203**, 6
- Park, D., Woo, J.-H., Denney, K. D., & Shin, J. 2013, *ApJ*, **770**, 87
- Park, D., Woo, J.-H., Treu, T., et al. 2012b, *ApJ*, **747**, 30
- Pei, L., Fausnaugh, M. M., Barth, A. J., et al. 2017, arXiv:1702.01177
- Peterson, B. M. 1993, *PASP*, **105**, 247
- Peterson, B. M. 2014, *SSRv*, **183**, 253
- Peterson, B. M., Bentz, M. C., Desroches, L.-B., et al. 2005, *ApJ*, **632**, 799
- Peterson, B. M., Ferrarese, L., Gilbert, K. M., et al. 2004, *ApJ*, **613**, 682
- Peterson, B. M., McHardy, I. M., Wilkes, B. J., et al. 2000, *ApJ*, **542**, 161
- Peterson, B. M., Pogge, R. W., Wanders, I., Smith, S. M., & Romanishin, W. 1995, *PASP*, **107**, 579
- Proga, D., & Kallman, T. R. 2004, *ApJ*, **616**, 688
- Proga, D., & Kurosawa, R. 2010, in ASP Conf. Ser. 427, Accretion and Ejection in AGN: A Global View, ed. L. Maraschi et al. (San Francisco, CA: ASP), 41
- Rodríguez-Pascual, P. M., Alloin, D., Clavel, J., et al. 1997, *ApJS*, **110**, 9
- Ruan, J. J., Anderson, S. F., Cales, S. L., et al. 2016, *ApJ*, **826**, 188
- Runco, J. N., Cosens, M., Bennert, V. N., et al. 2016, *ApJ*, **821**, 33
- Runnoe, J. C., Cales, S., Ruan, J. J., et al. 2016, *MNRAS*, **455**, 1691
- Rybicki, G. B., & Kleyana, J. T. 1994, in ASP Conf. Ser. 69, Reverberation Mapping of the Broad-Line Region in Active Galactic Nuclei, ed. P. M. Gondhalekar, K. Horne, & B. M. Peterson (San Francisco, CA: ASP), 85
- Sakata, Y., Minezaki, T., Yoshii, Y., et al. 2010, *ApJ*, **711**, 461
- Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, **737**, 103
- Schmidt, M. 1963, *Natur*, **197**, 1040
- Sergeev, S. G., Doroshenko, V. T., Golubinskiy, Y. V., Merkulova, N. I., & Sergeeva, E. A. 2005, *ApJ*, **622**, 129

- Shapovalova, A. I., Popović, L. Č., Burenkov, A. N., et al. 2010, *A&A*, **517**, A42
- Shappee, B. J., Prieto, J. L., Grupe, D., et al. 2014, *ApJ*, **788**, 48
- Shemmer, O., Uttley, P., Netzer, H., & McHardy, I. M. 2003, *MNRAS*, **343**, 1341
- Simpson, C. 2005, *MNRAS*, **360**, 565
- Skjelboe, A., Pancoast, A., Treu, T., et al. 2015, *MNRAS*, **454**, 144
- Suganuma, M., Yoshii, Y., Kobayashi, Y., et al. 2006, *ApJ*, **639**, 46
- Tody, D. 1986, *Proc. SPIE*, **627**, 733
- Trakhtenbrot, B., & Netzer, H. 2012, *MNRAS*, **427**, 3081
- Tully, R. B., Shaya, E. J., Karachentsev, I. D., et al. 2008, *ApJ*, **676**, 184
- Turner, T. J., Miller, L., Reeves, J. N., & Braito, V. 2016, arXiv:1602.05158
- Ulrich, M.-H., & Horne, K. 1996, *MNRAS*, **283**, 748
- Valenti, S., Sand, D. J., Barth, A. J., et al. 2015, *ApJL*, **813**, L36
- van der Walt, S., Colbert, S. C., & Varoquaux, G. 2011, *Computing in Science & Engineering*, **13**, 22
- van Dokkum, P. G. 2001, *PASP*, **113**, 1420
- van Groningen, E., & Wanders, I. 1992, *PASP*, **104**, 700
- Vaughan, S., Edelson, R., Warwick, R. S., & Uttley, P. 2003, *MNRAS*, **345**, 1271
- Vazquez, B., Galianni, P., Richmond, M., et al. 2015, *ApJ*, **801**, 127
- Vestergaard, M., & Peterson, B. M. 2006, *ApJ*, **641**, 689
- Wandel, A., Peterson, B. M., & Malkan, M. A. 1999, *ApJ*, **526**, 579
- Wanders, I., Peterson, B. M., Alloin, D., et al. 1997, *ApJS*, **113**, 69
- Weedman, D. W. 1977, *ARA&A*, **15**, 69
- Welsh, W. F. 1999, *PASP*, **111**, 1347
- Whittle, M. 1992, *ApJS*, **79**, 49
- Willite, B. C., Vanden Berk, D. E., Kron, R. G., et al. 2005, *ApJ*, **633**, 638
- Wills, B. J., & Browne, I. W. A. 1986, *ApJ*, **302**, 56
- Woo, J.-H., Yoon, Y., Park, S., Park, D., & Kim, S. C. 2015, *ApJ*, **801**, 38
- Yuan, F., & Narayan, R. 2014, *ARA&A*, **52**, 529
- Zu, Y., Kochanek, C. S., Kozłowski, S., & Udalski, A. 2013, *ApJ*, **765**, 106
- Zu, Y., Kochanek, C. S., & Peterson, B. M. 2011, *ApJ*, **735**, 80